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*ASTR 2401*

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# Filters

Observational Astronomy

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# Agenda

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- ❖ Weather for Tonight
- ❖ Semester Project Proposal
- ❖ Lab / Observing Notebook Due Friday 09 / 30 by 4:25 PM
  - ❖ Cutoff TBD
- ❖ Lecture:
  - ❖ Problems with Images
  - ❖ Filters

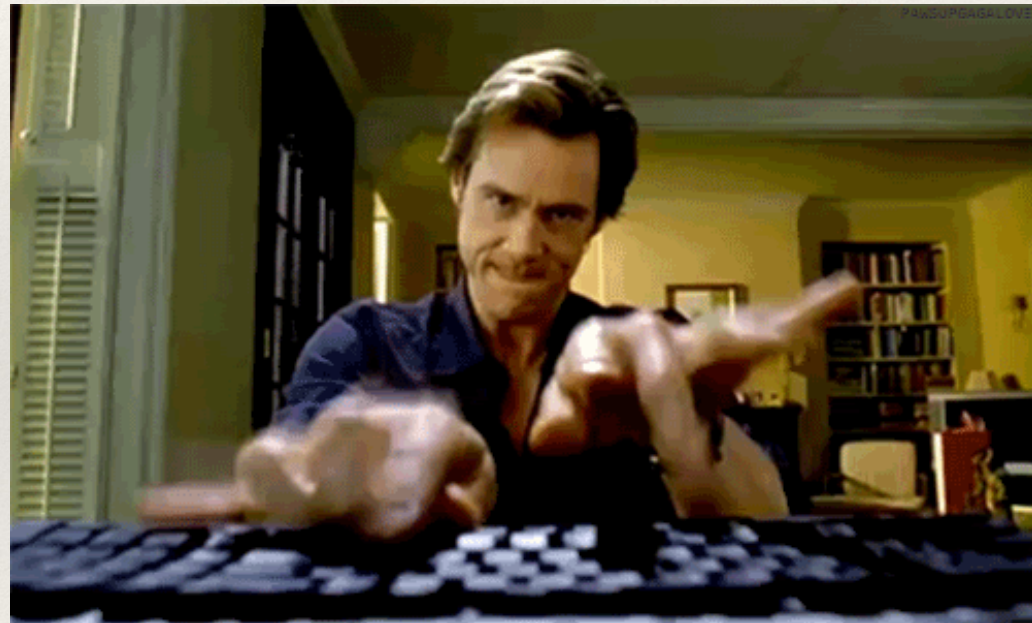


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# Semester Project Proposal

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- ❖ 2-page double spaced (12pt font 1-inch margins)
- ❖ Must contain:
  - ❖ Scientific background
  - ❖ Observational techniques to be used
  - ❖ Targets to be observed
  - ❖ Equipment (telescopes, CCDs, filters, etc.)
  - ❖ Plan for data reduction/analysis
  - ❖ At least two peer-review references cited
- ❖ Due in class October 4th





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# Problems with your Data

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So what will the data look like when things go wrong?

Defects that are always present:

- Bad Pixels
- Cosmic Rays
- Bleed Trails

More serious issues

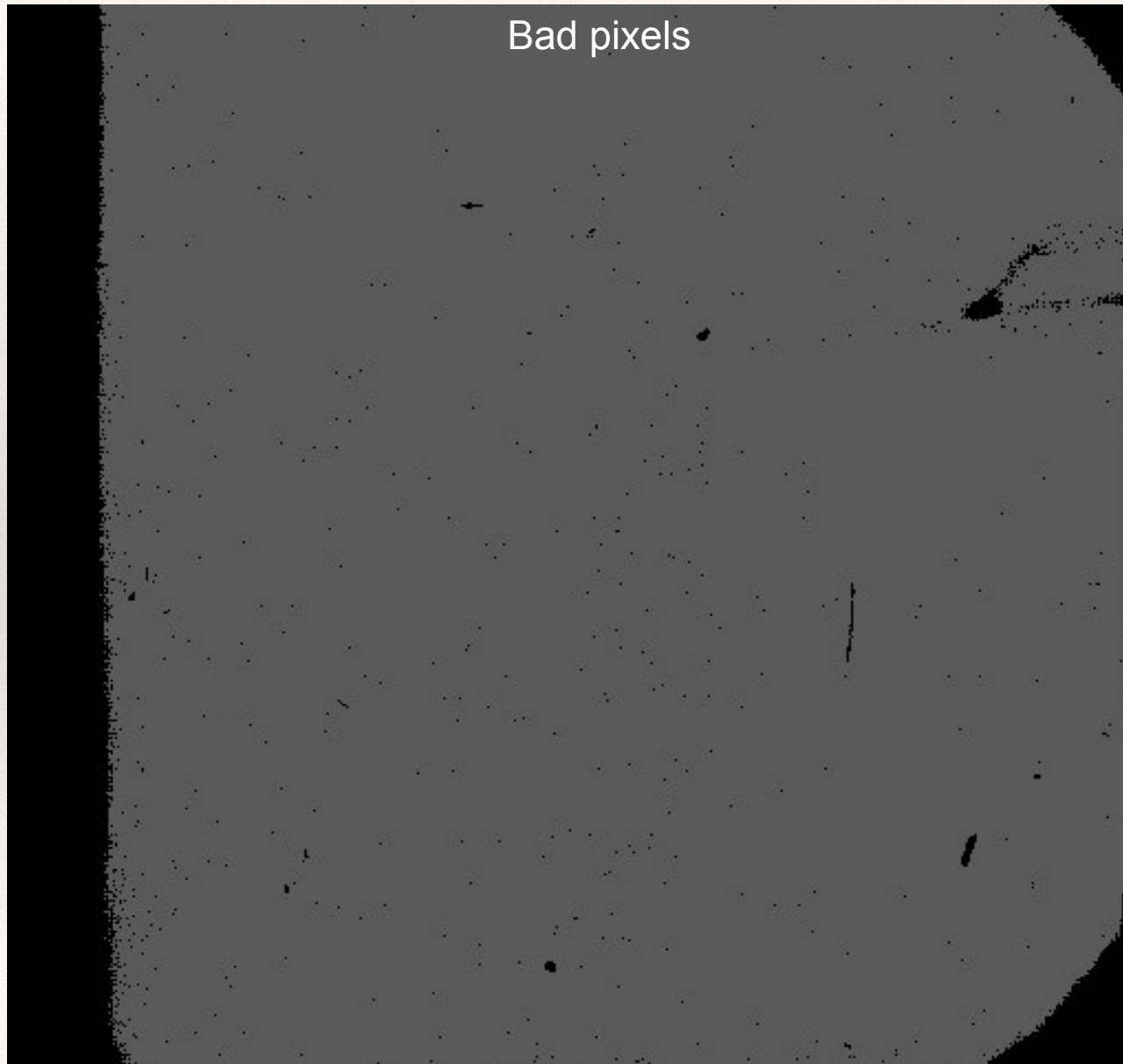
- Saturated Image
- Vignetting (sometimes unavoidable)
- Out of Focus
- Bad Tracking
- Scattered light



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# Problems with your Data

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# Problems with your Data

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Bleed Trails



Credit: Kurtis Williams

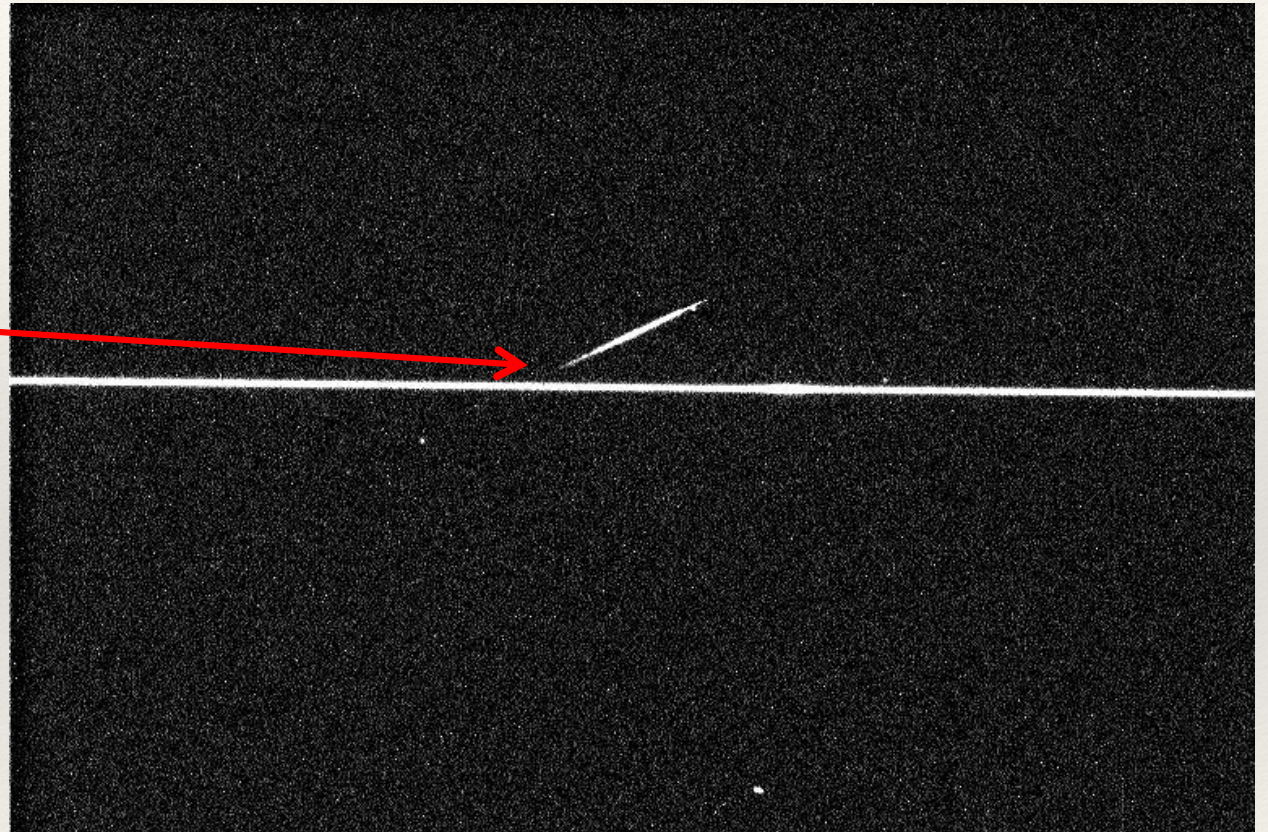


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# Problems with your Data

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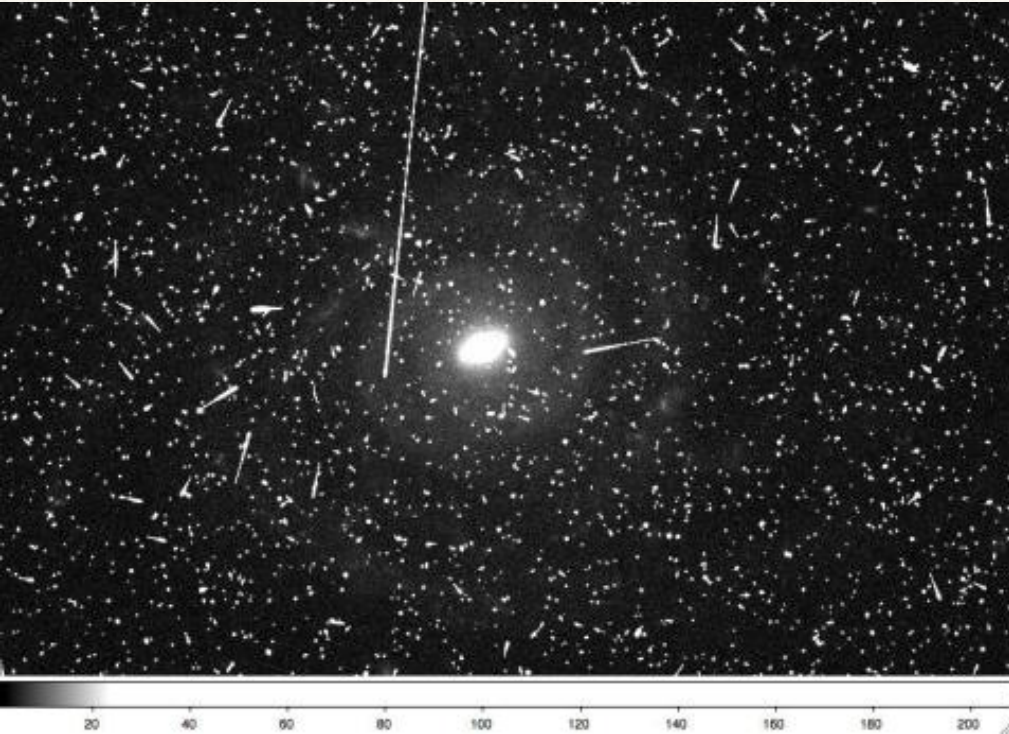
Cosmic Rays



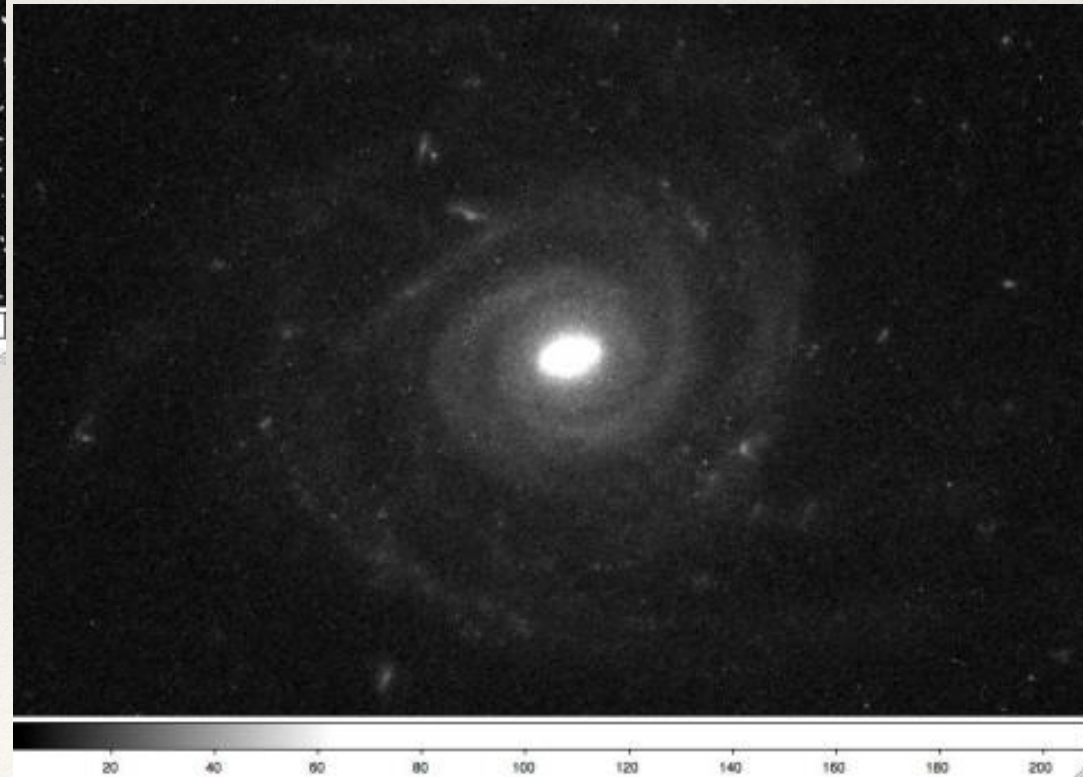


# Problems with your Data

Cosmic Rays



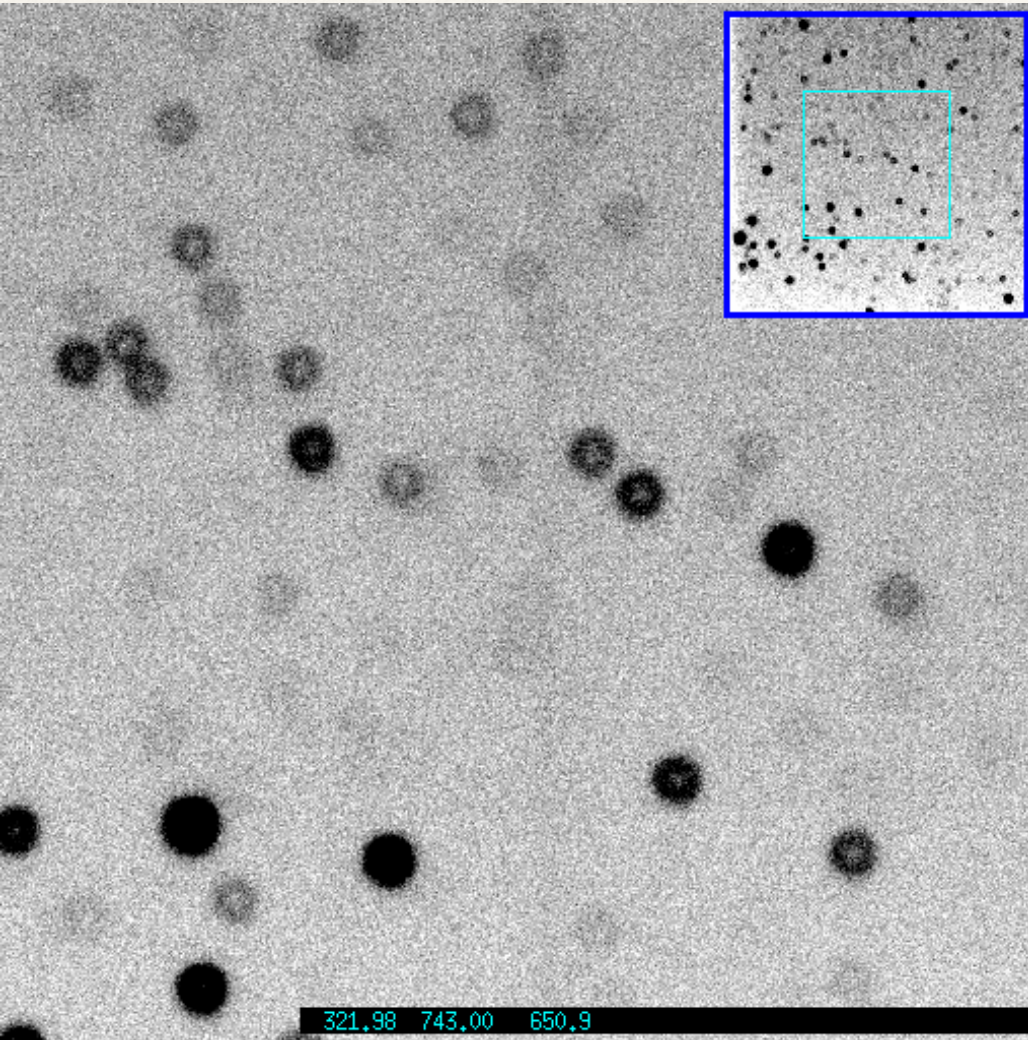
Courtesy: <http://blog.galaxyzoo.org/2010/04/12/how-to-handle-hubble-images/>



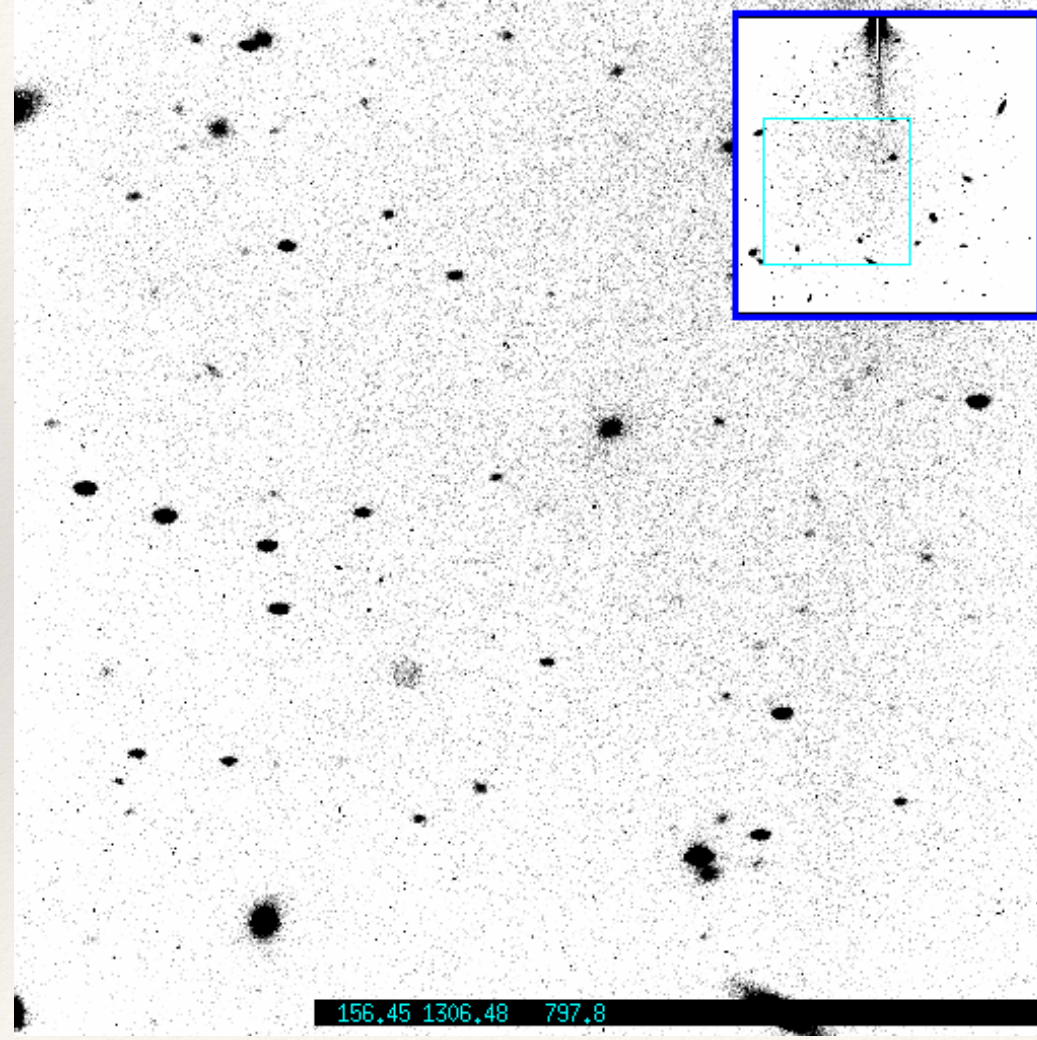


# Problems with your Data

Out of focus



Bad tracking



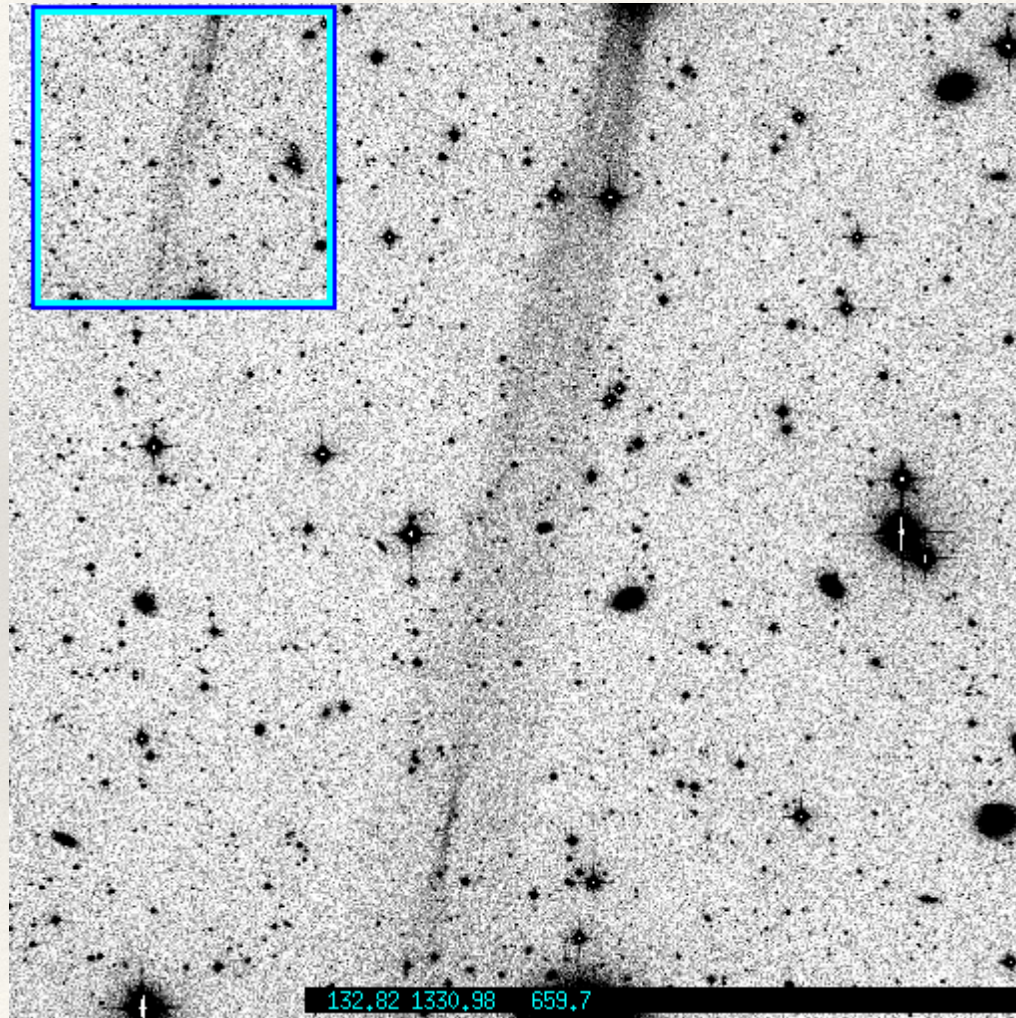


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# Problems with your Data

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Scattered Light



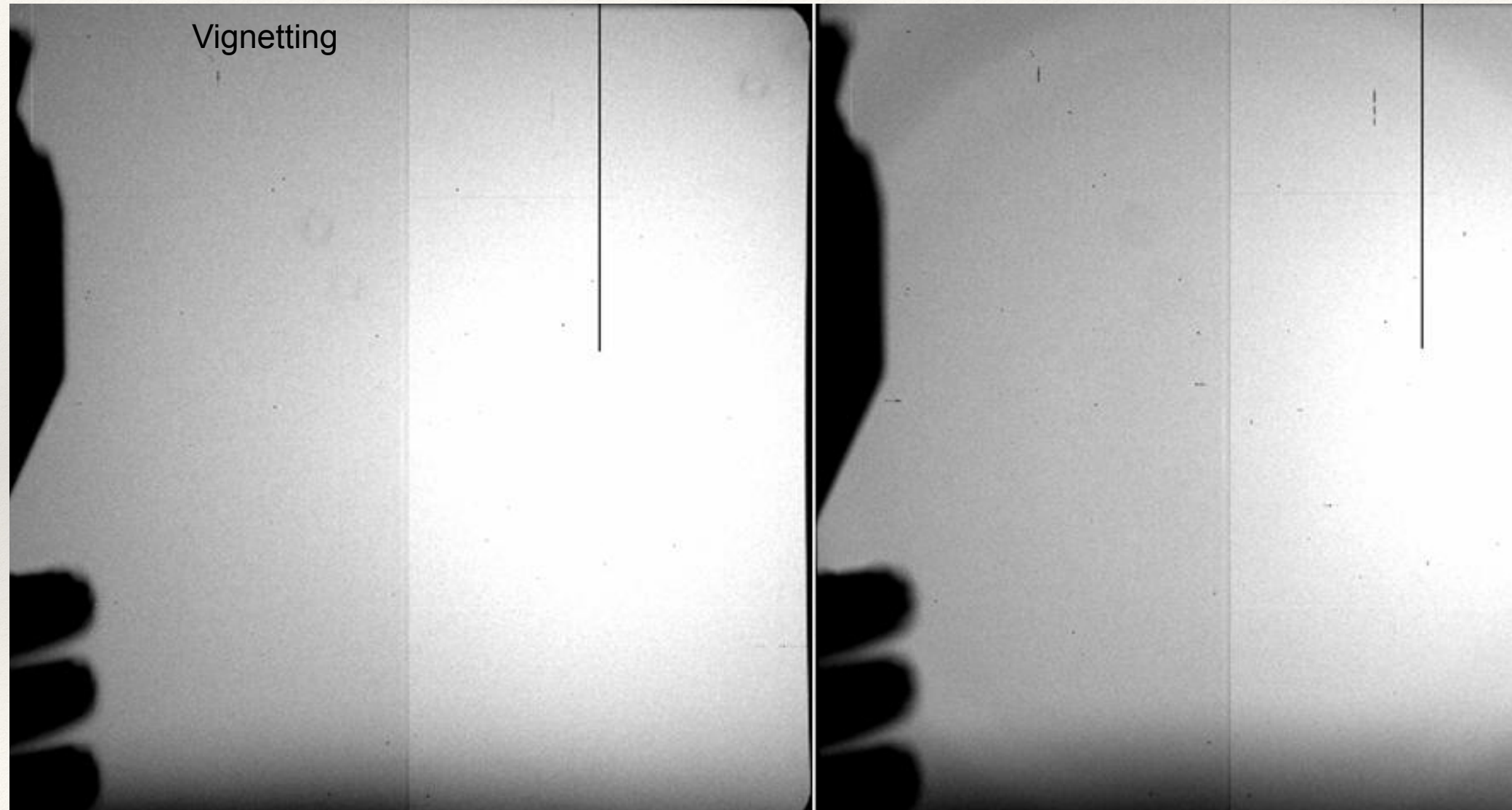


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# Problems with your Data

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Vignetting





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# Filters

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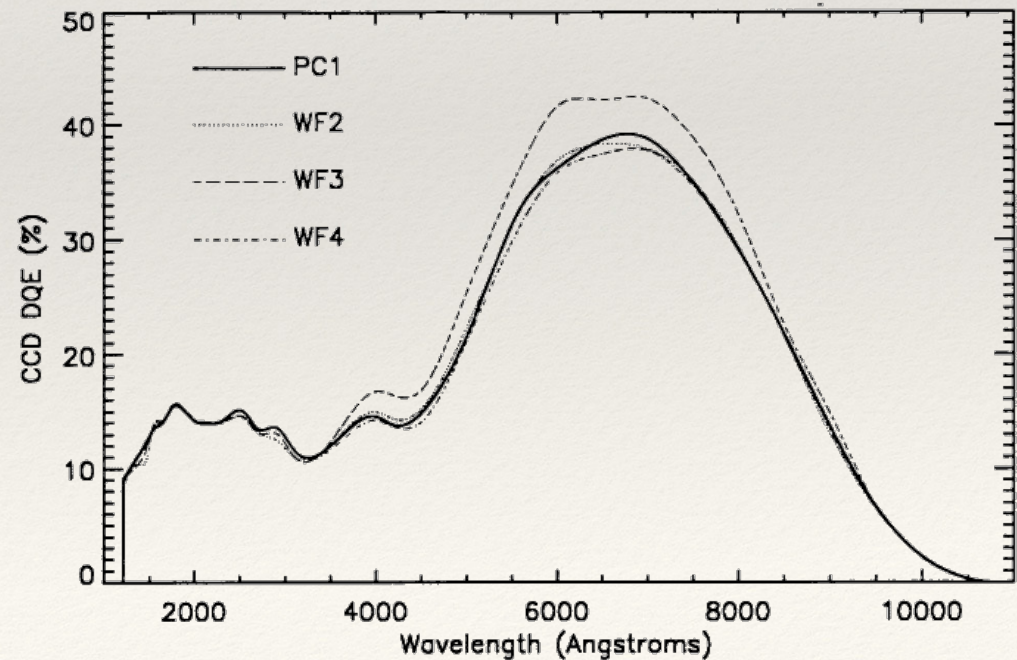
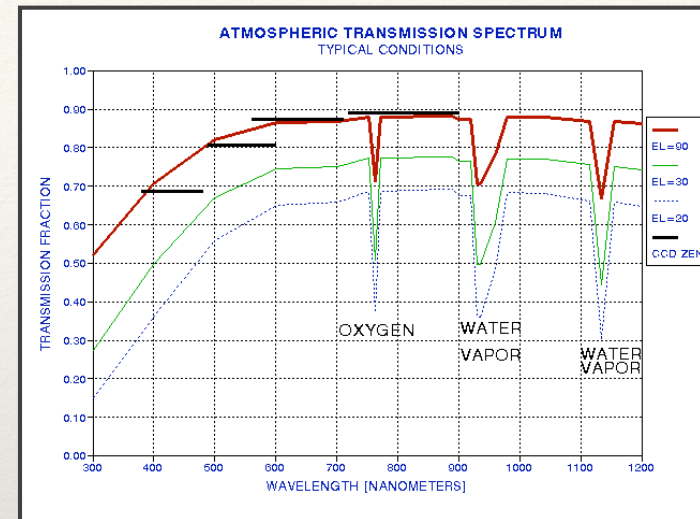
# Filters

Filters are optical elements designed to transmit only certain wavelengths of light.

They are defined by their **central wavelength** and **bandwidth** (how wide a range of wavelengths the transmit).

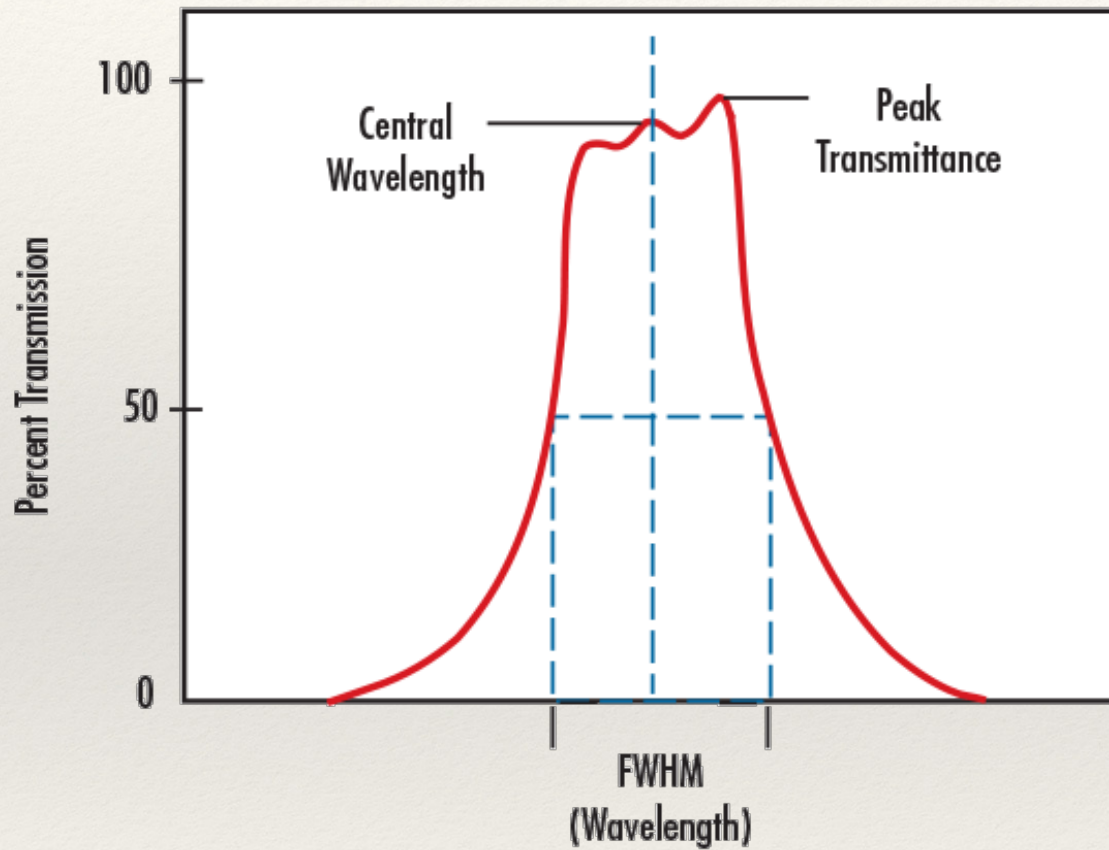
Without a filter, all light that passes through telescope hits detector.

Amount of light then determined by QE of detector, and transmission of optics and the atmosphere.





# Filters



Bandwidth is usually defined as the full width at half maximum (FWHM)

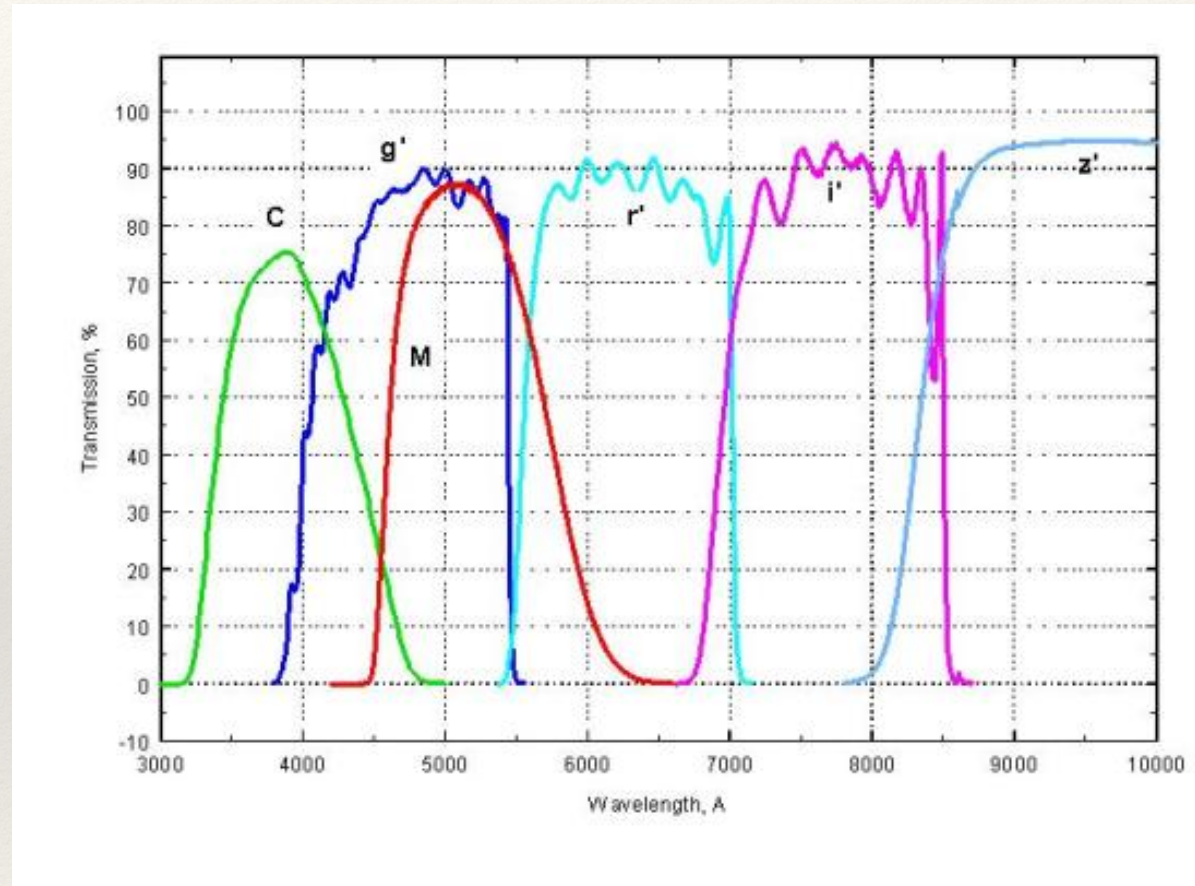


# Filters

The point of filters is to allow you to observe only certain wavelengths of light at a time.

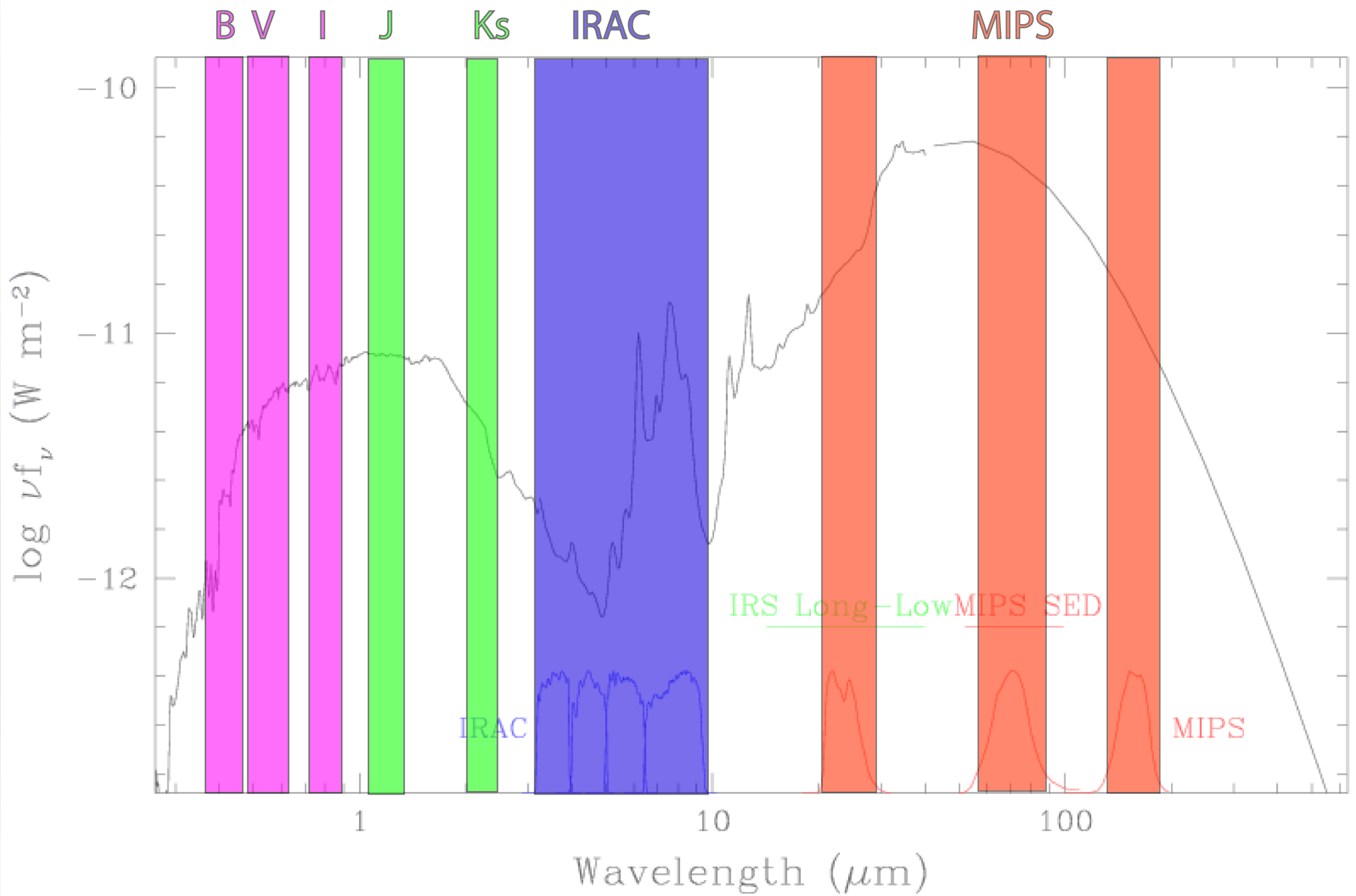
Provides color information

Enables analysis of physical properties of objects.



[www.noao.edu/kpno/mosaic/filters](http://www.noao.edu/kpno/mosaic/filters)

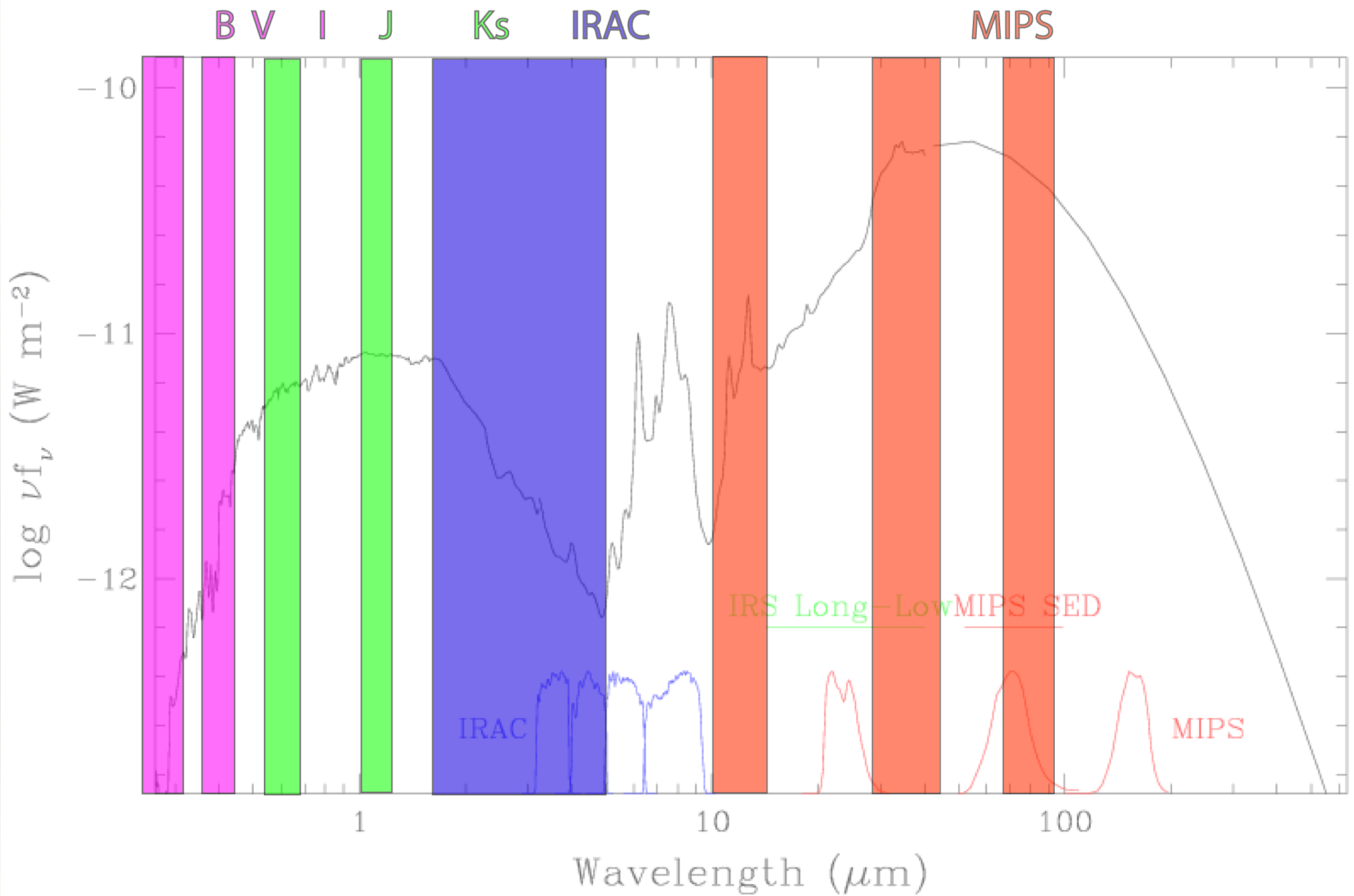




$z=0$

M82 (<http://sings.stsci.edu>)

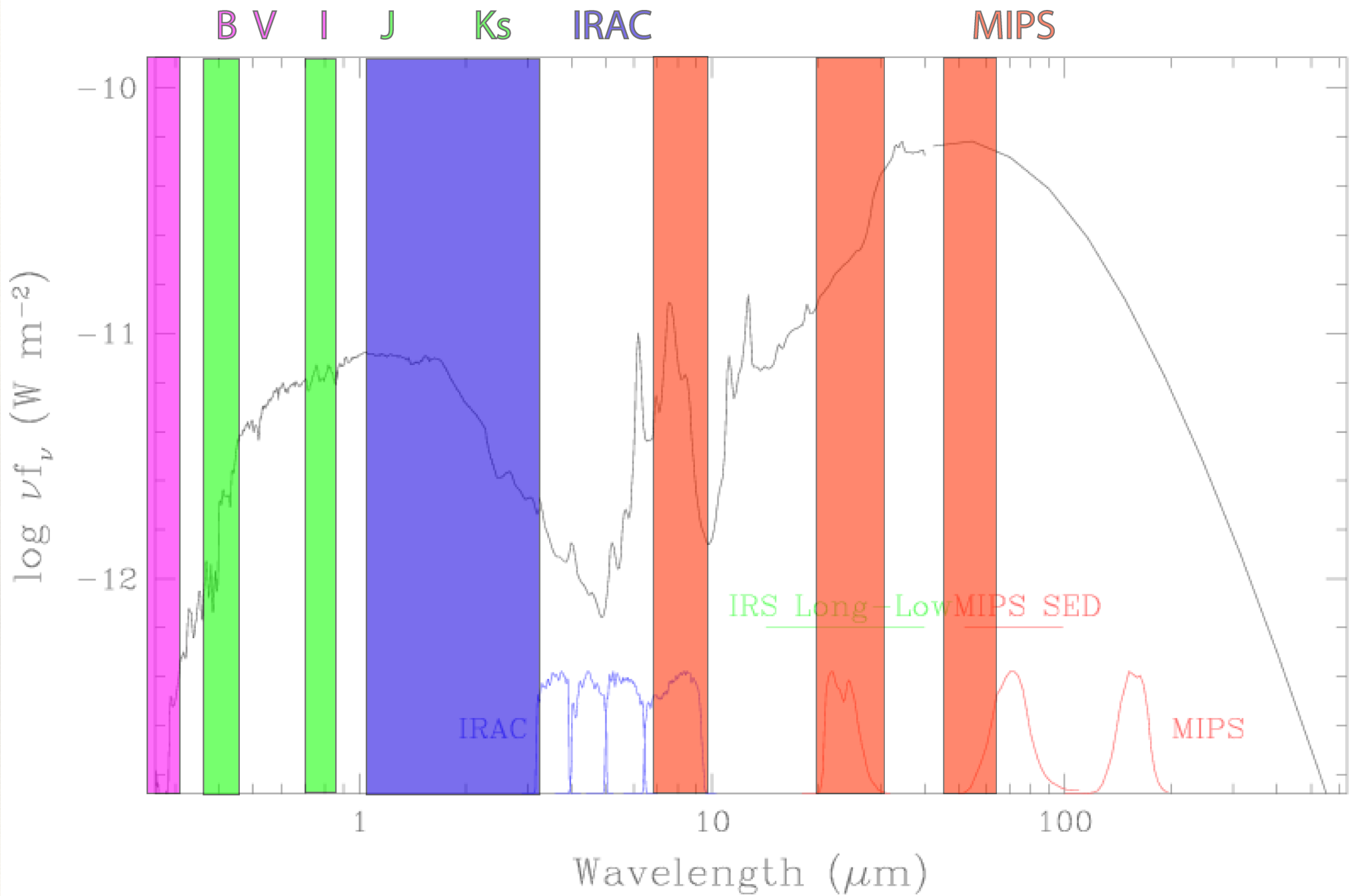




$z=1$

M82 (<http://sings.stsci.edu>)





$z=2$

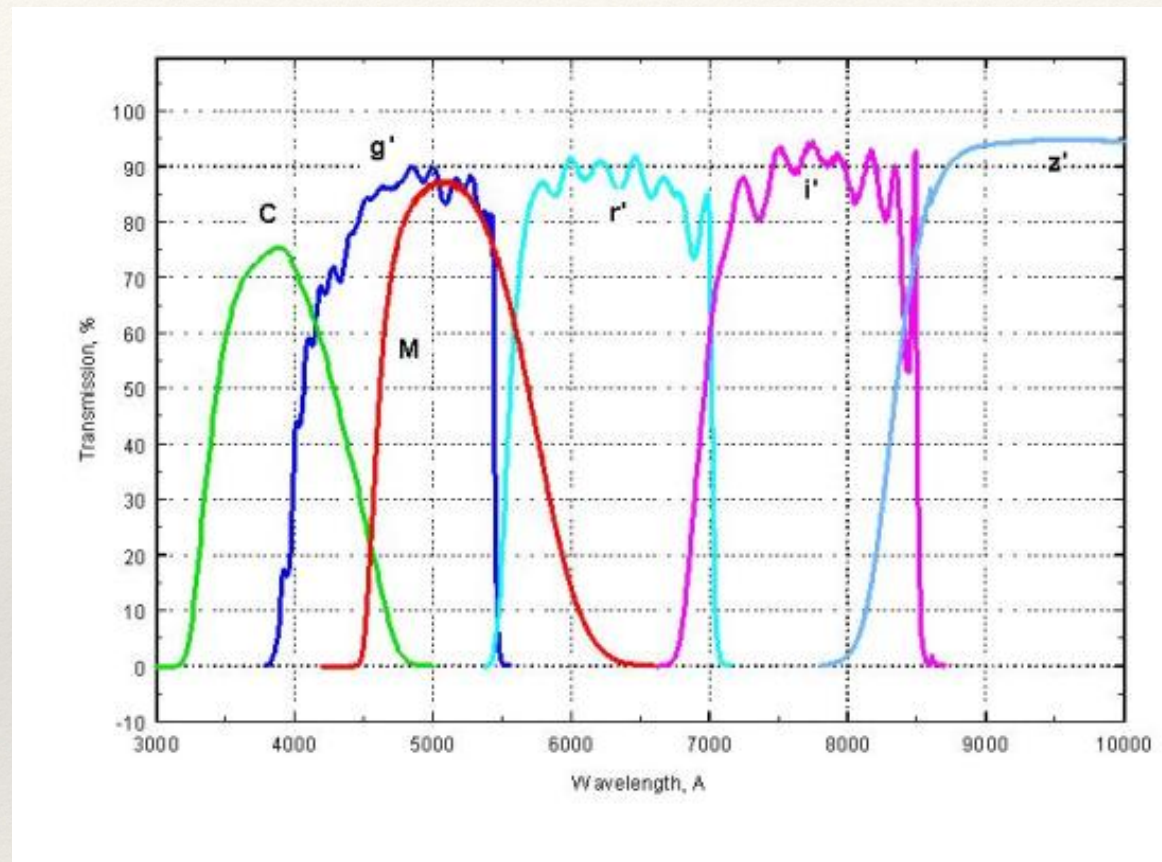
M82 (<http://sings.stsci.edu>)



# Filters

There are a host of different filters used in astronomy, typically defined based upon some combination of the physics being studied and atmospheric transmission.

- ❖ For all filters one can define a *resolution*,
  - $R = \lambda / \Delta\lambda$
- ❖ Filters are generally considered “broad”, “medium”, or “narrowband” depending upon the resolution.
- ❖ Very Rough Guide
  - ❖ Broad:  $R < 10$
  - ❖ Medium:  $R \sim 10\text{-}50$
  - ❖ Narrow:  $R > 50$

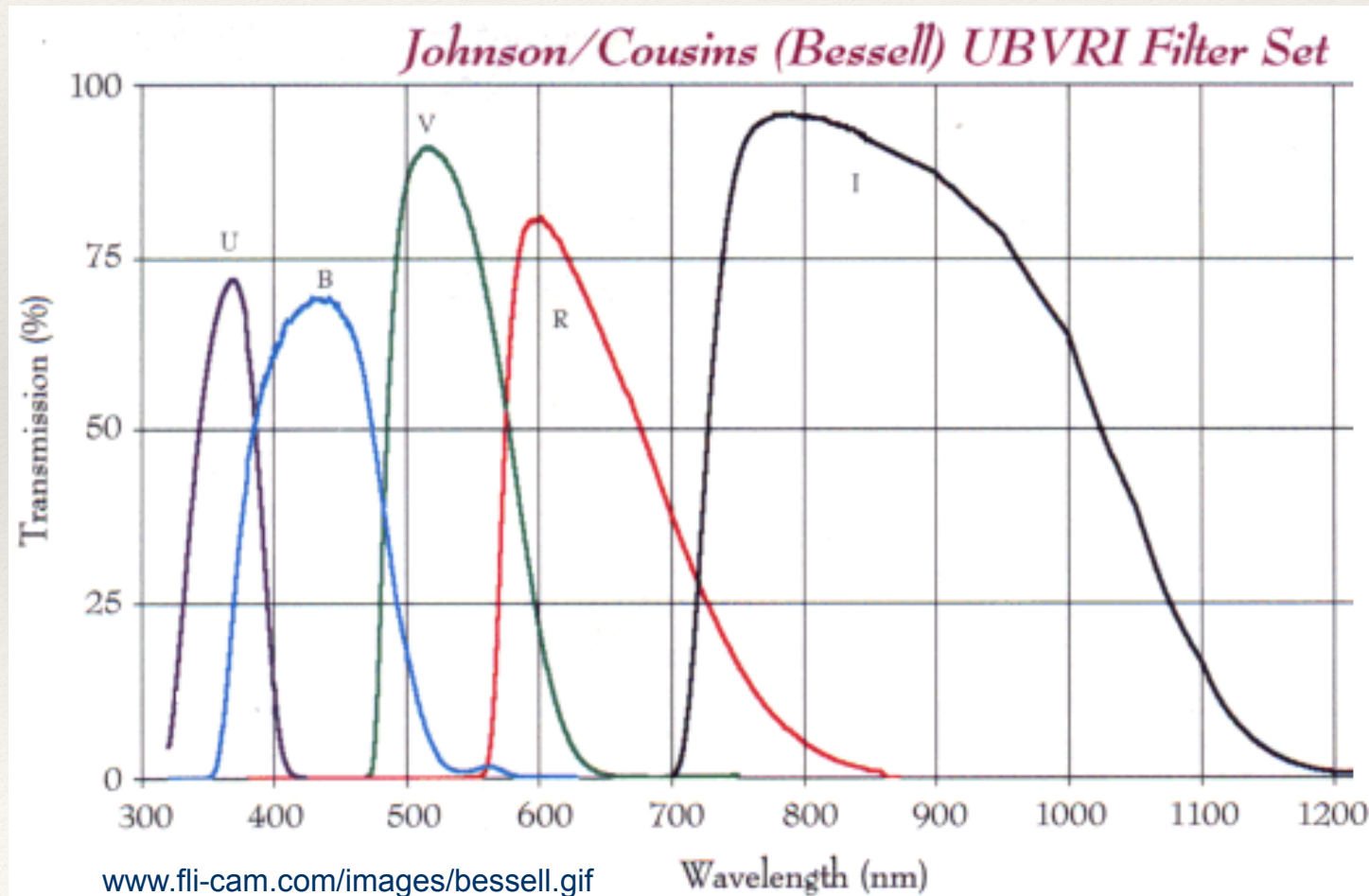




# Broadband Filters

## Johnson-Cousins Filter Set

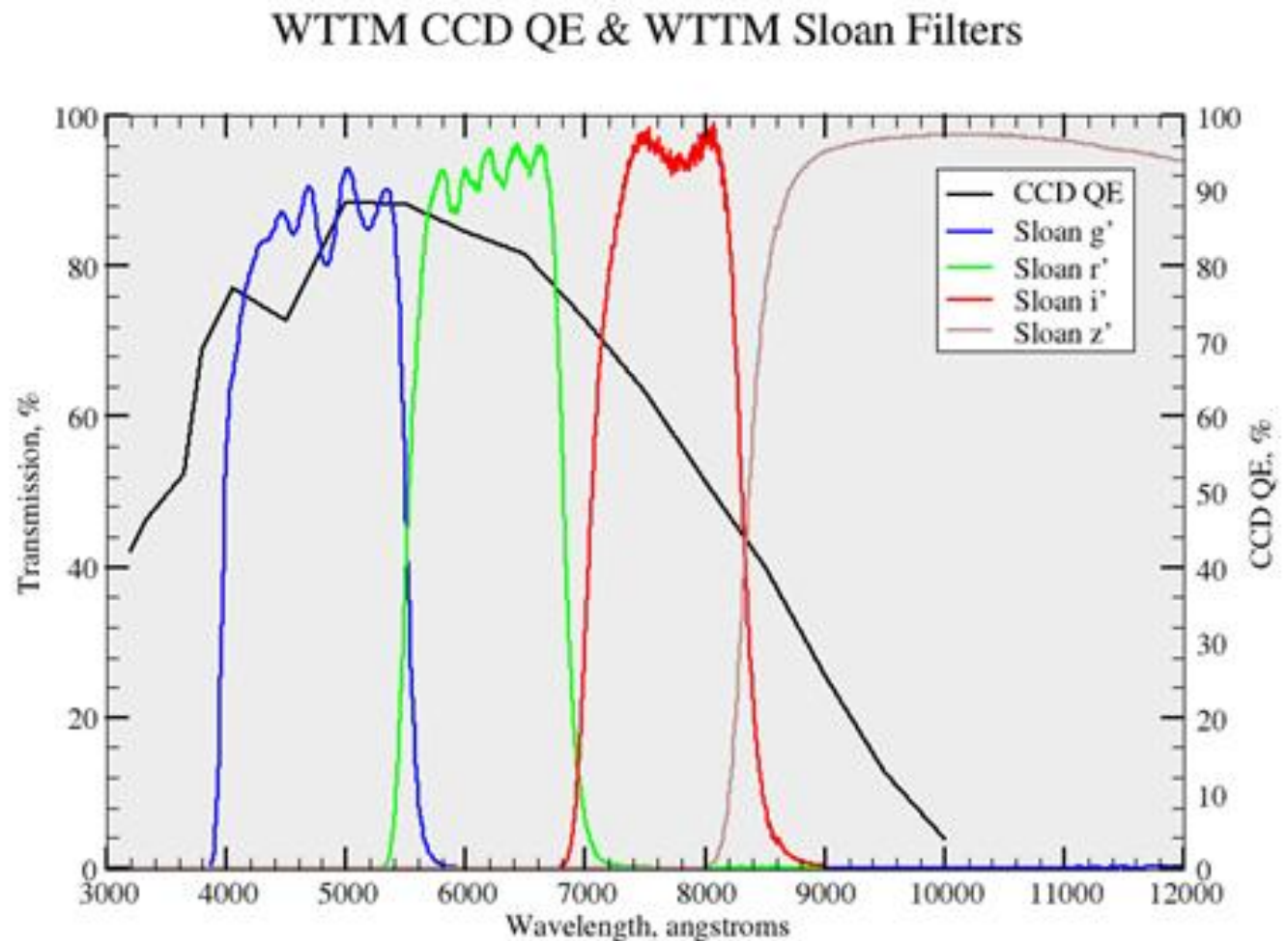
- ❖ UBVRI
- ❖ Classical filter set
  - ❖ Ultraviolet (U)
  - ❖ Blue (B)
  - ❖ Visible (V)
  - ❖ Red (R)
  - ❖ “Infrared” (I)
- ❖ Optical and very near IR
- ❖ Wavelength cutoffs not very sharp



# Broadband Filters

## Sloan Filter Set

- ❖  $g'r'i'z'$
- ❖ Relatively new
- ❖ Sharp cutoffs
- ❖ Used for Sloan Digital Sky Survey
- ❖ Optical and very near IR
- ❖ No red cutoff for  $z'$





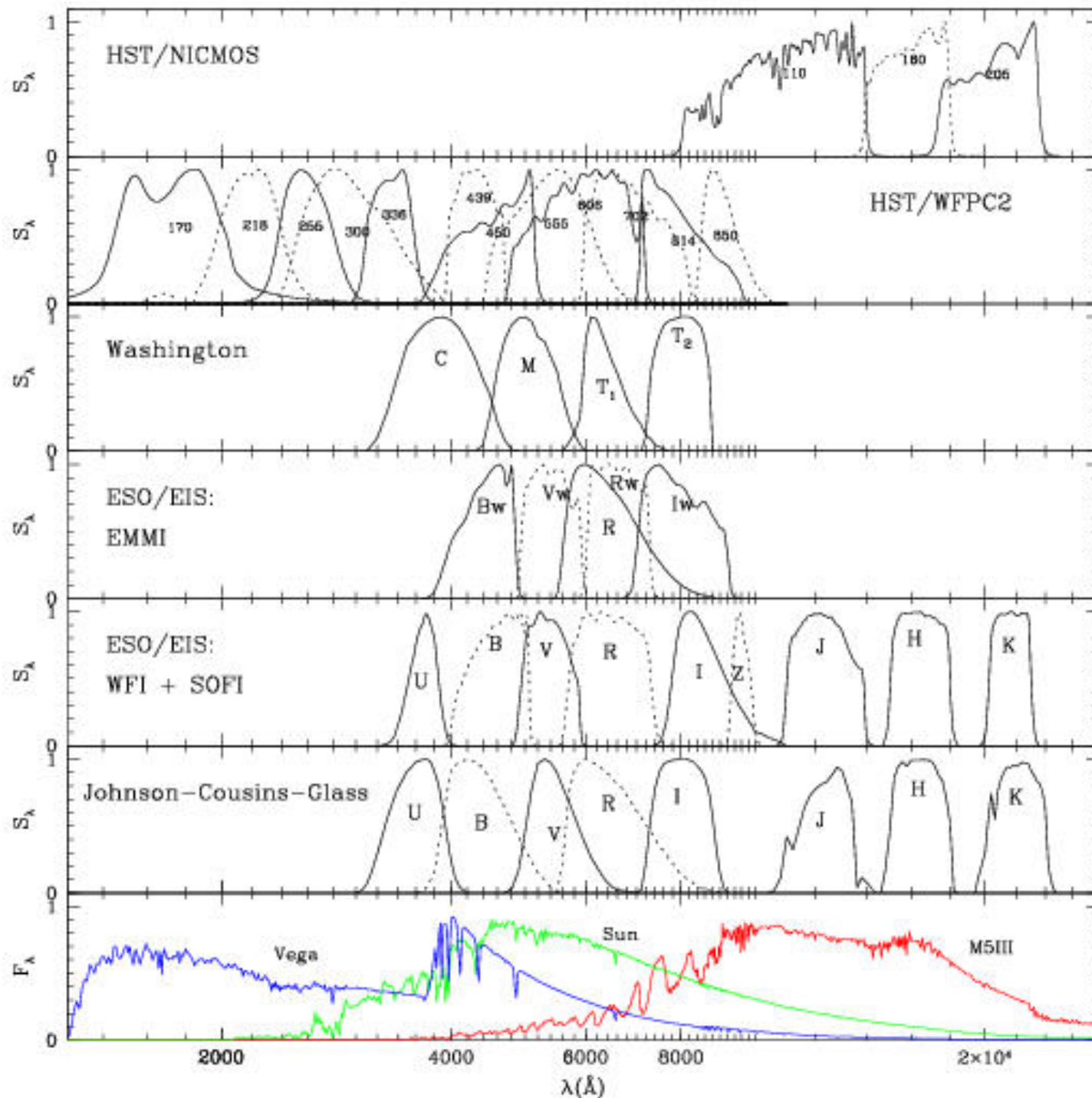


[http://wfc3.gsfc.nasa.gov/art2/gallery\\_images/filter-images/ir-wheel.jpg](http://wfc3.gsfc.nasa.gov/art2/gallery_images/filter-images/ir-wheel.jpg)

Many Different Systems

- Kron
- Washington
- ...

*Even filters on the same system will have slightly different transmission*





# eXtreme Deep Field (XDF) ■ *Hubble Space Telescope*

38135539

$z=8.7$



37796000

$z=8.5$



33436598

$z=8.6$



SN Primo

$z=1.55$



39546284

$z=10.3$

ACS/WFC F435W F606W

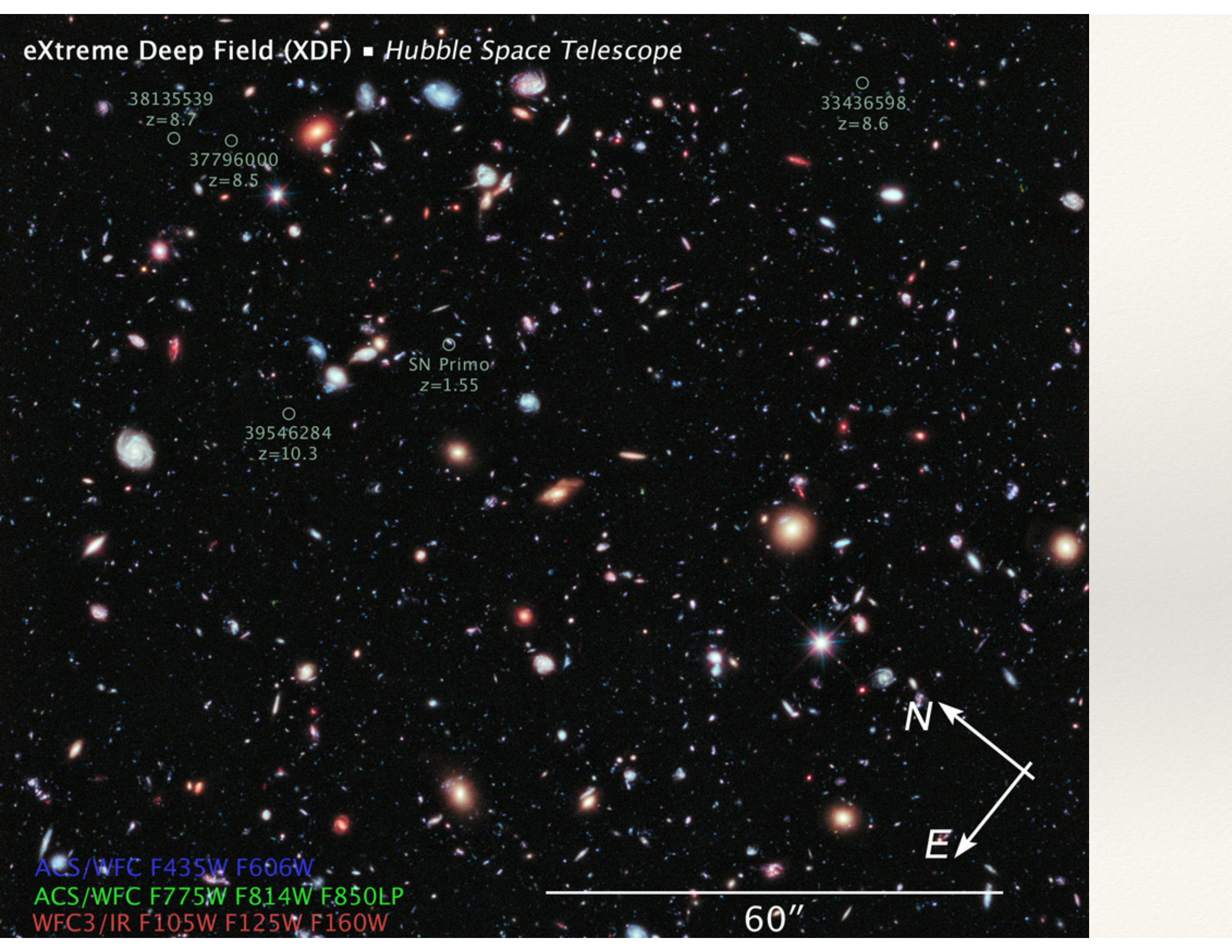
ACS/WFC F775W F814W F850LP

WFC3/IR F105W F125W F160W

N

E

60"

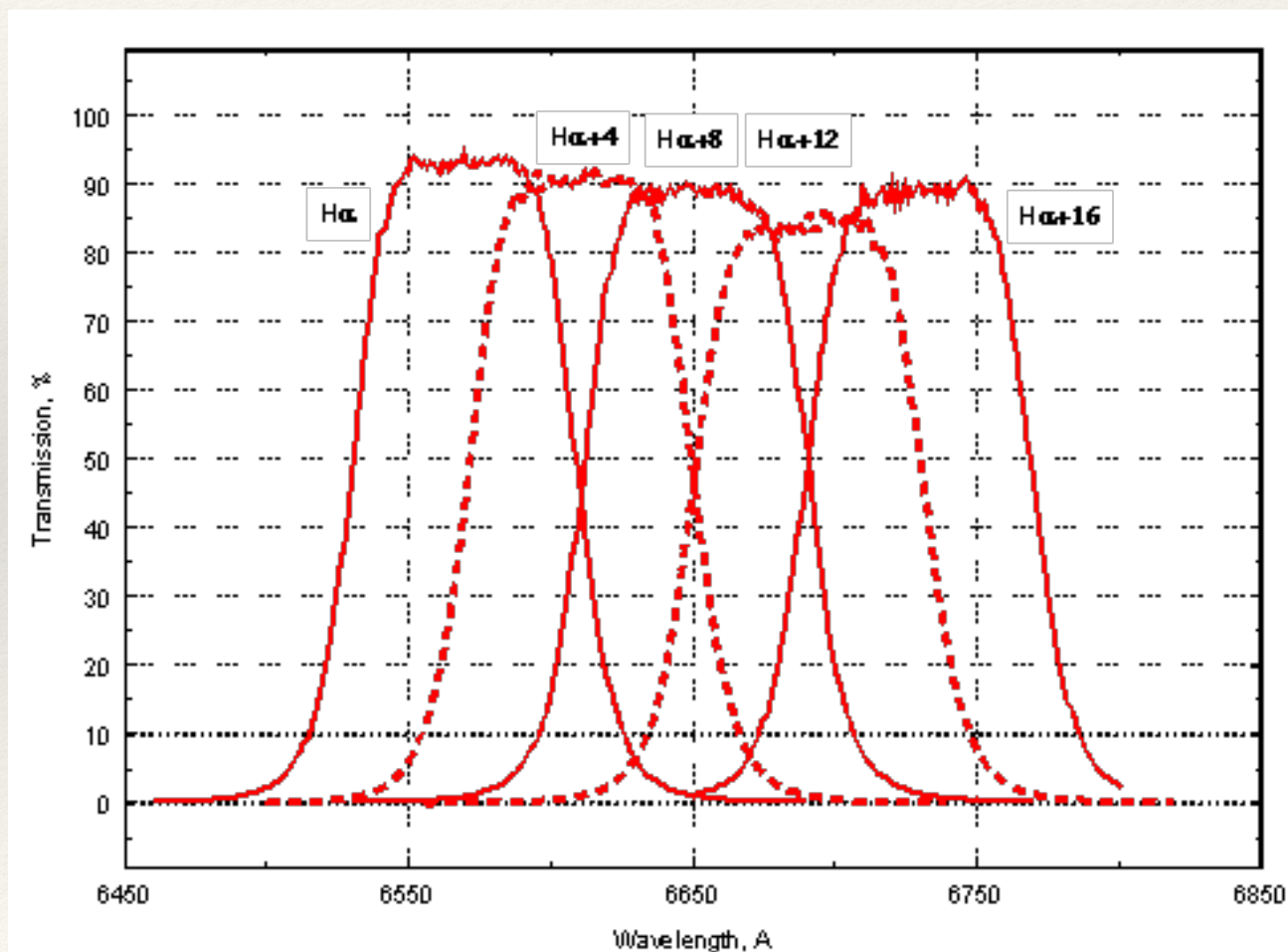




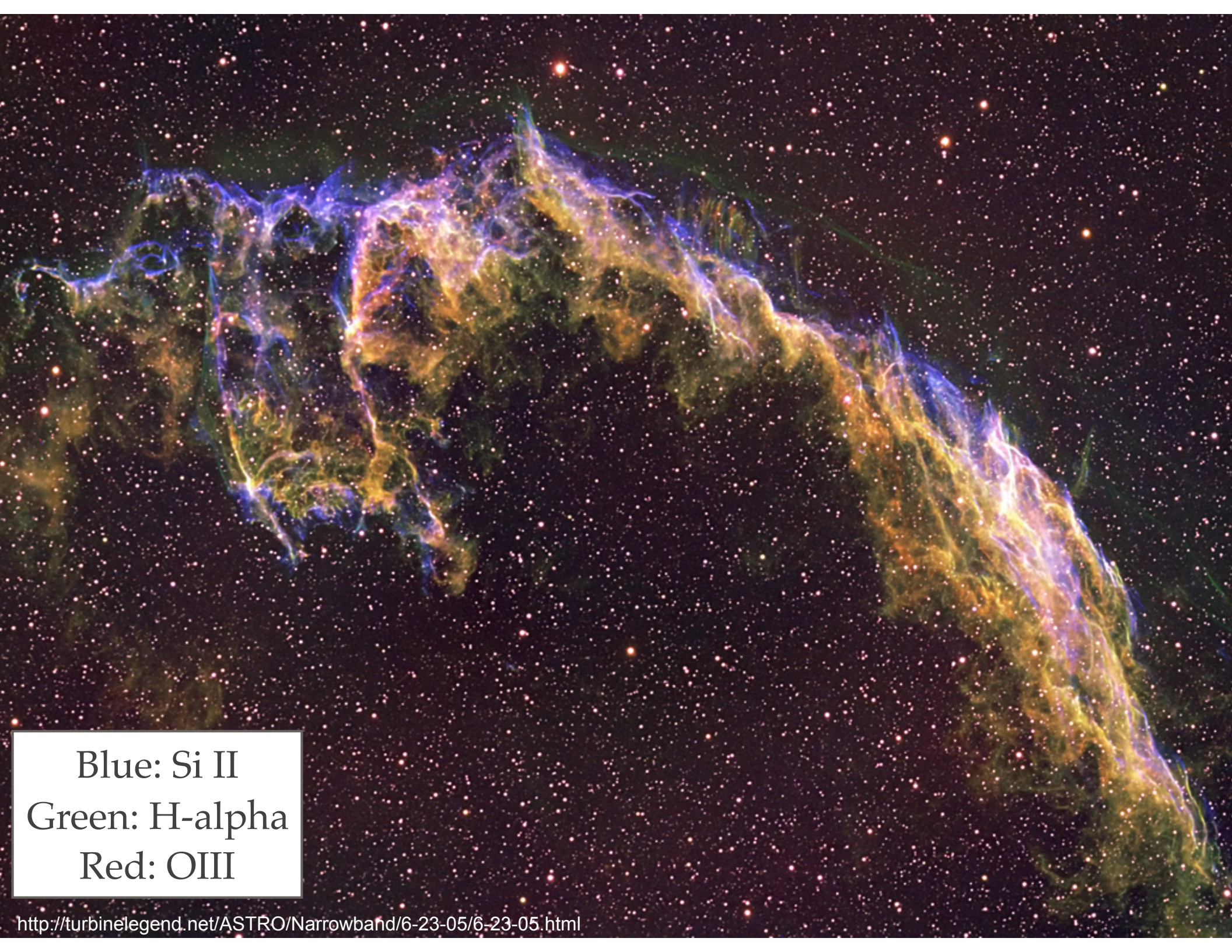
# Narrowband Filters

Typically target specific emission lines

- ❖ H-alpha
- ❖ Lyman-alpha
- ❖ OII
- ❖ OIII
- ❖ etc...





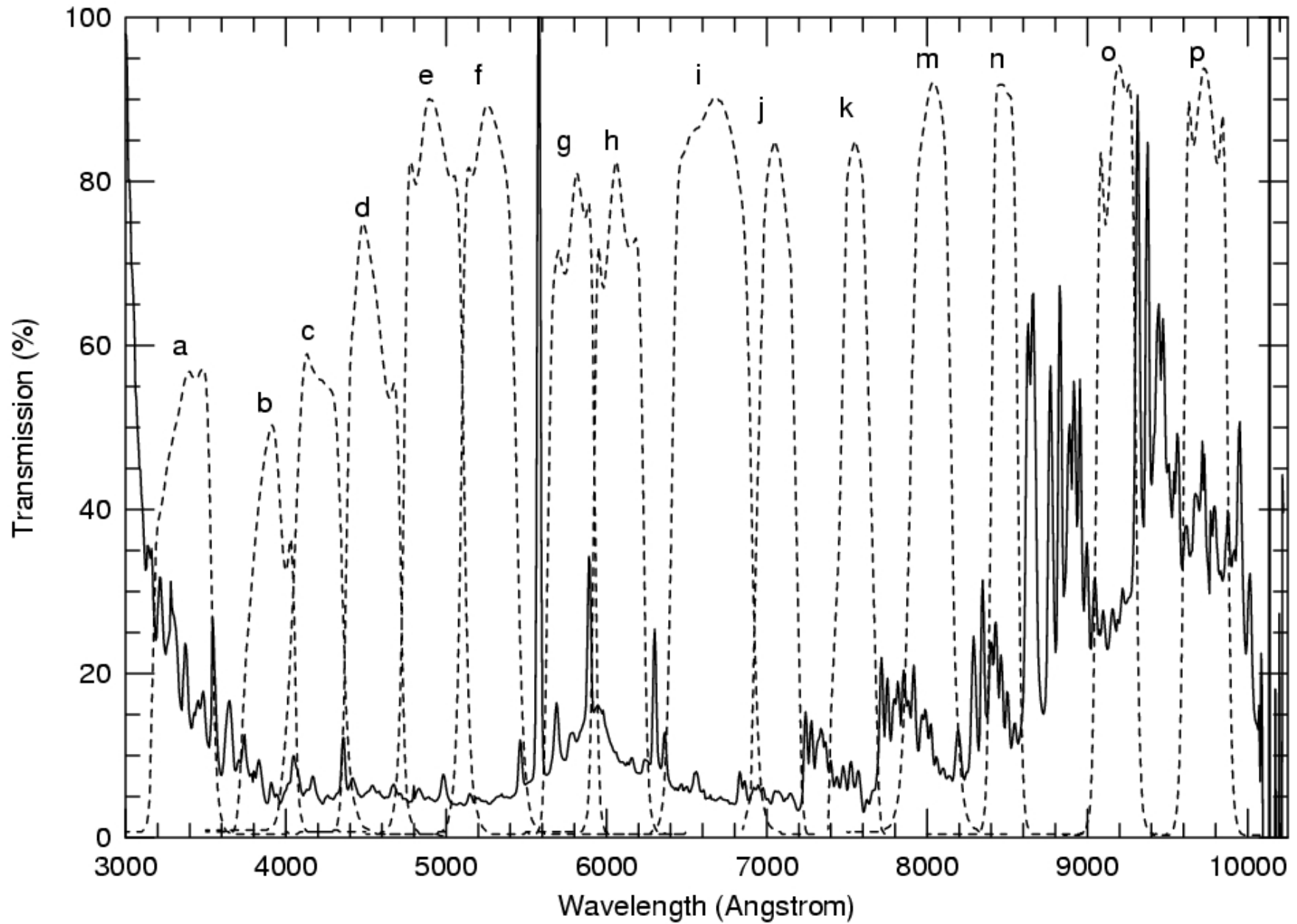


Blue: Si II

Green: H-alpha

Red: OIII

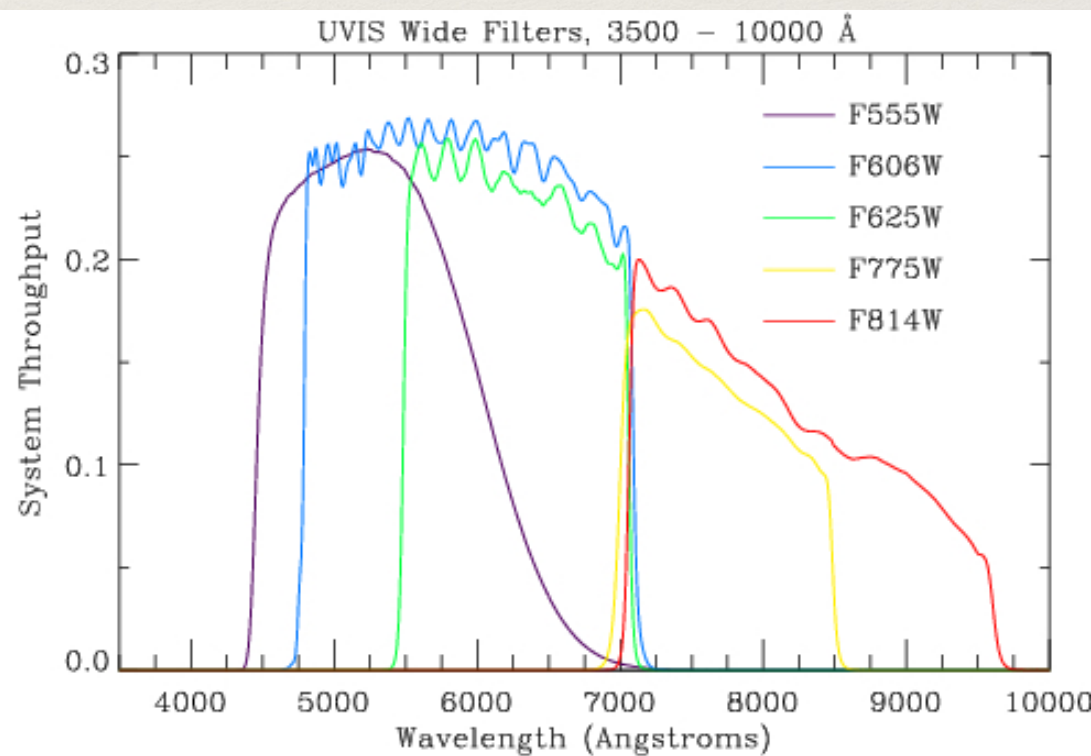
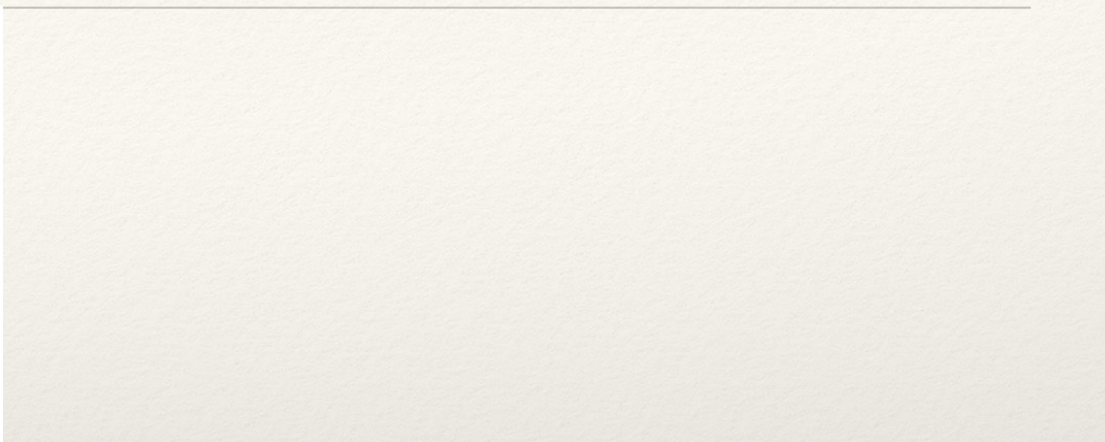
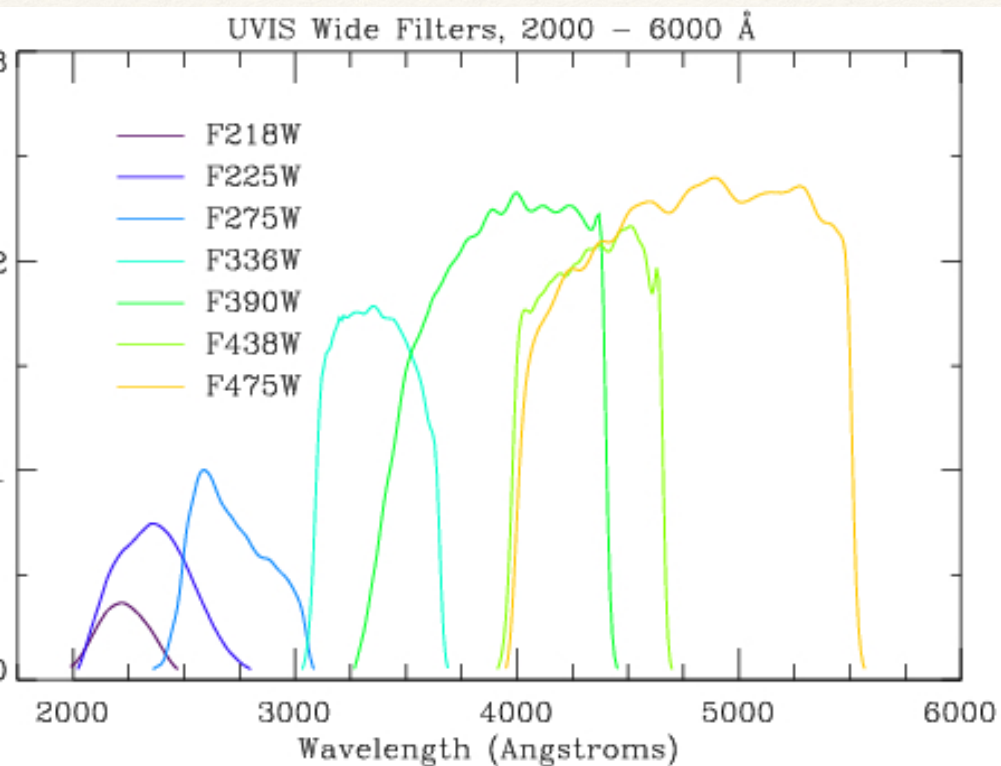




*Need a different filter for each redshift...*



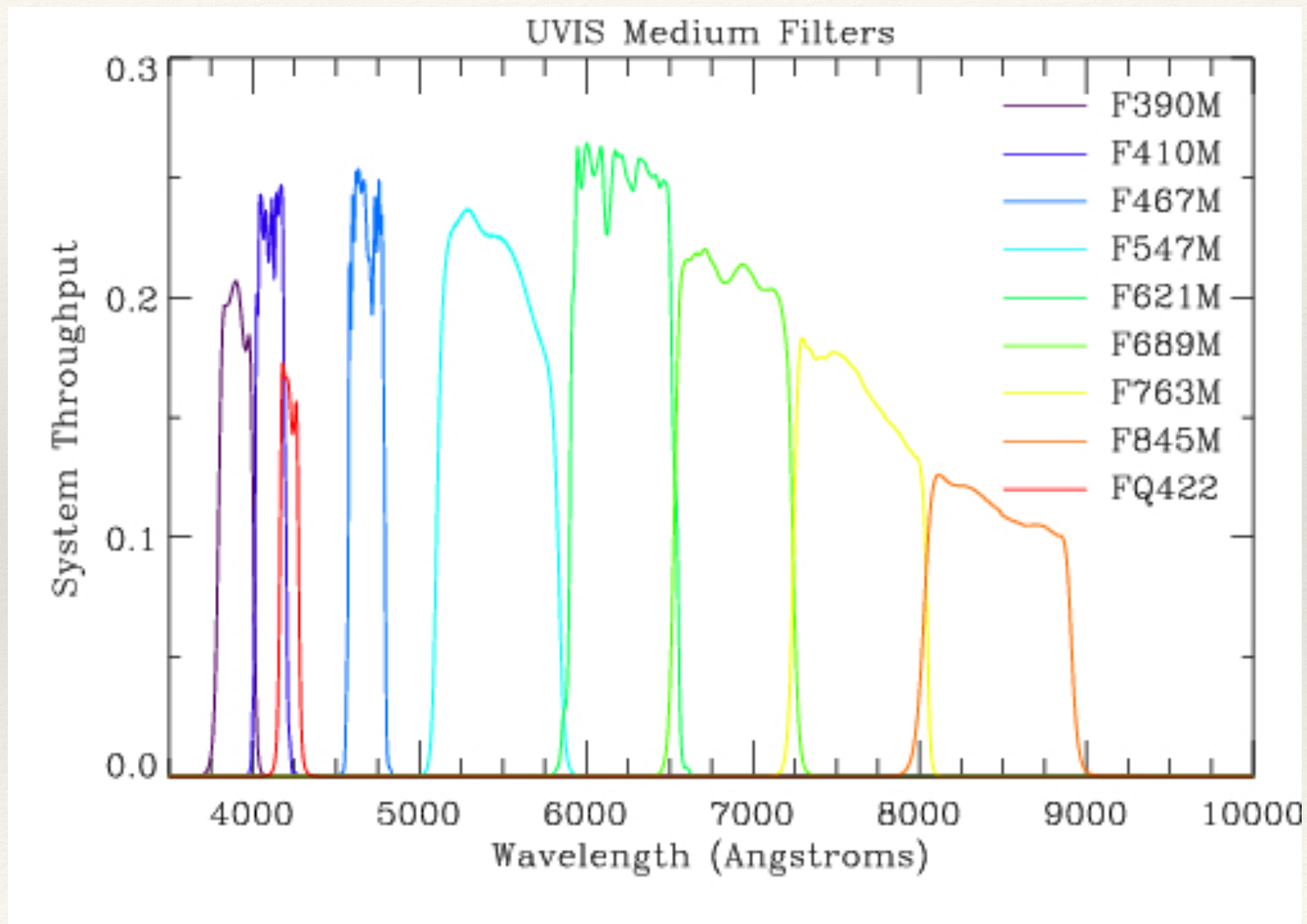
# WFC3 for HST



Broad band optical

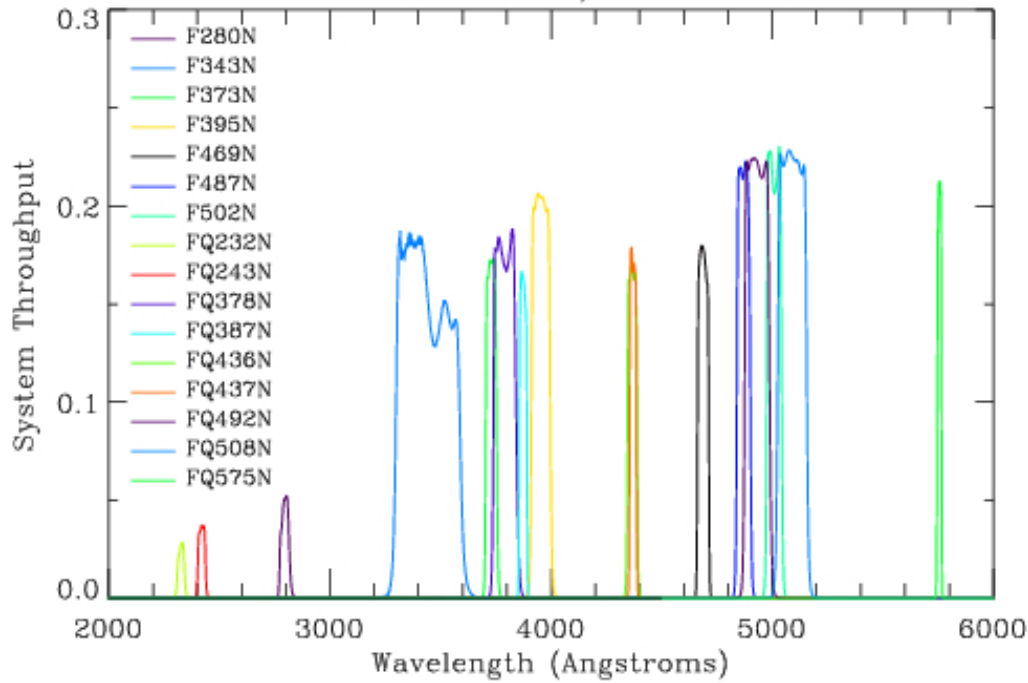
# WFC3 for HST

Medium band  
optical

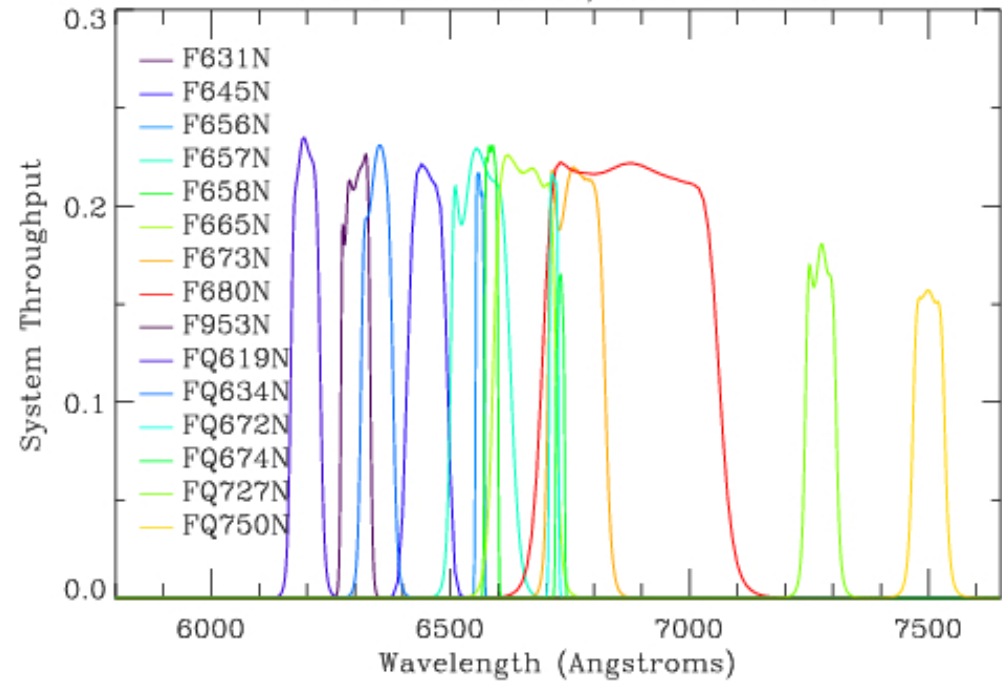




UVIS Narrow Filters, 2000 – 6000 Å

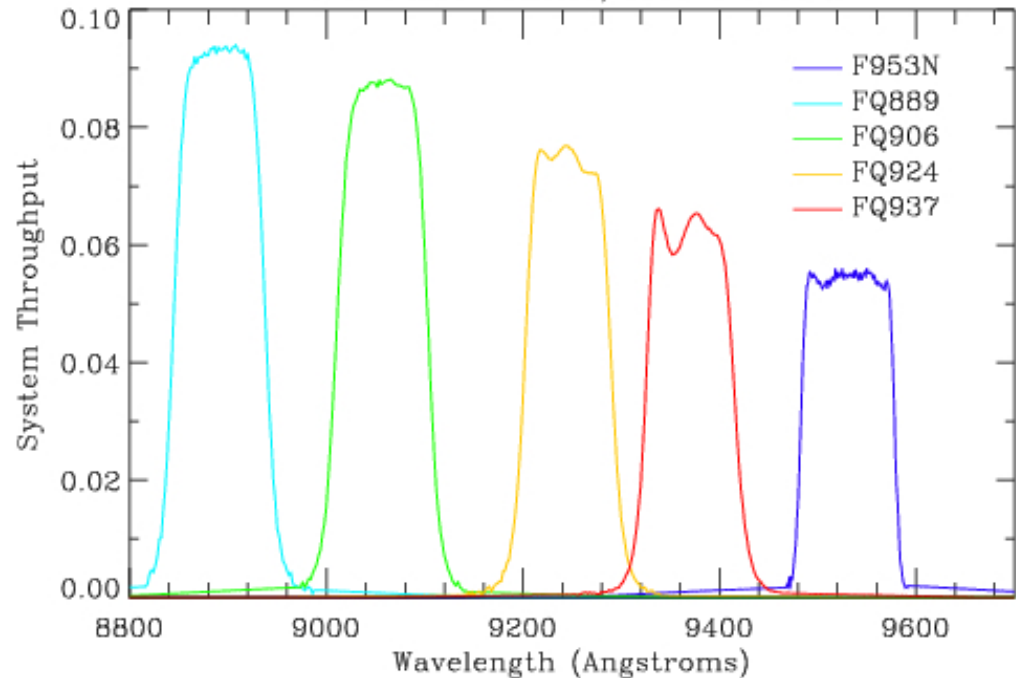


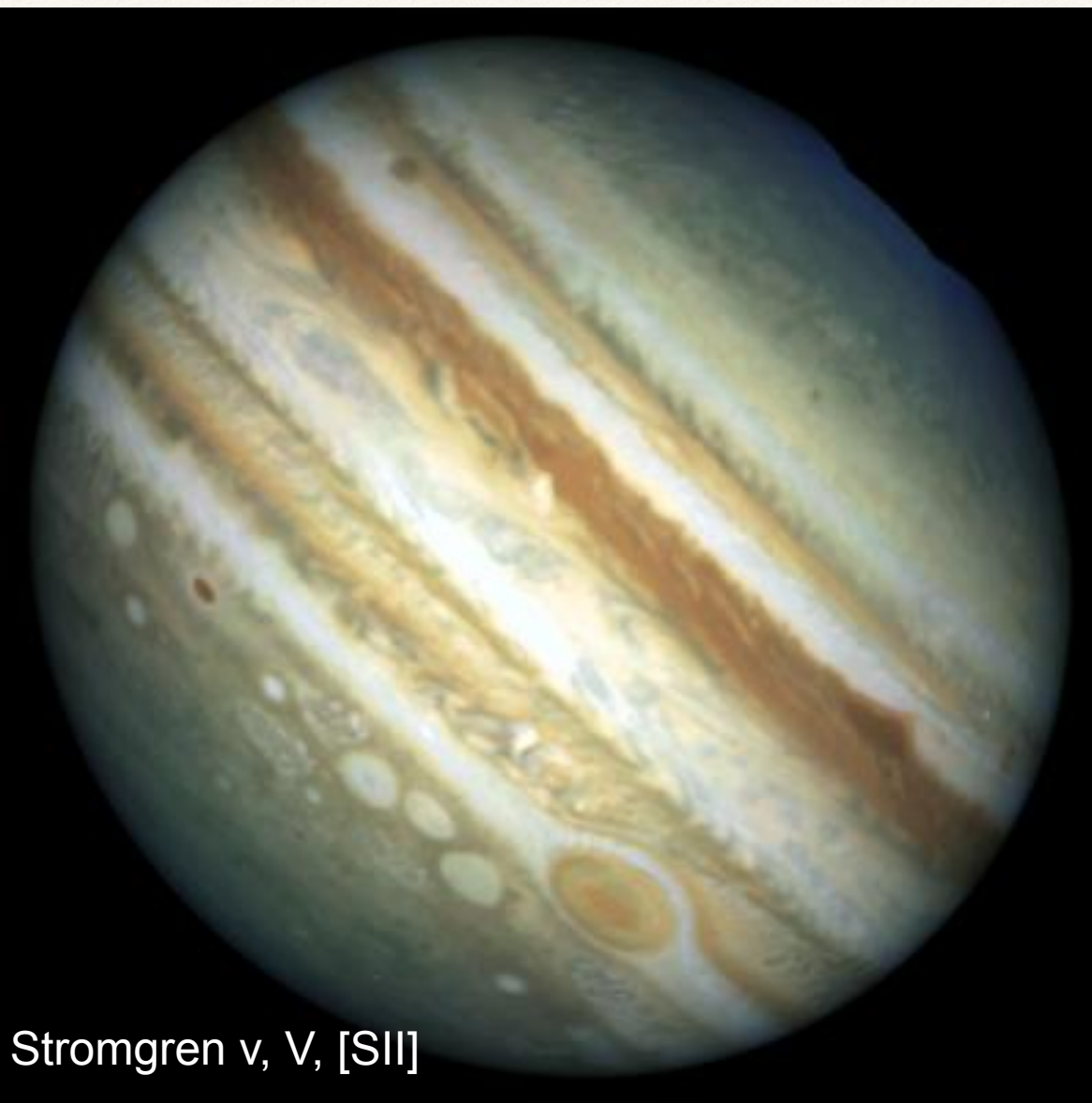
UVIS Narrow Filters, 6000 – 7650 Å



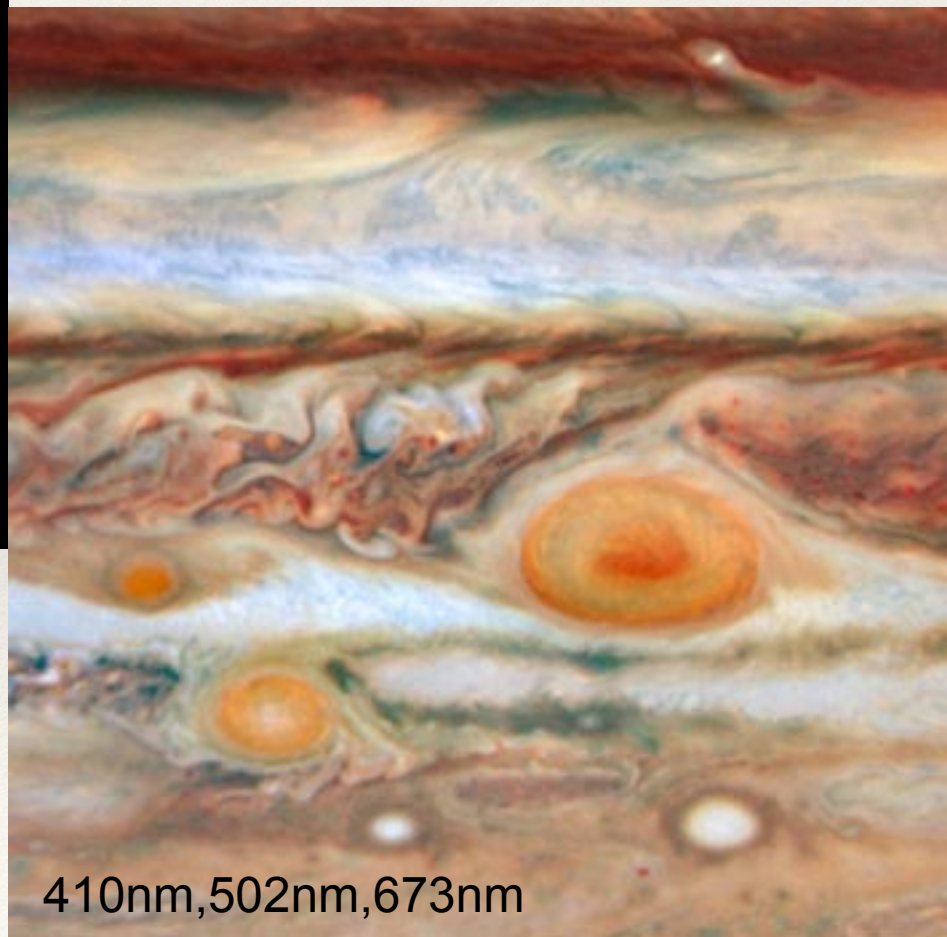
# WFC3 Narrow band optical

UVIS Narrow Filters, 8800 – 9700 Å





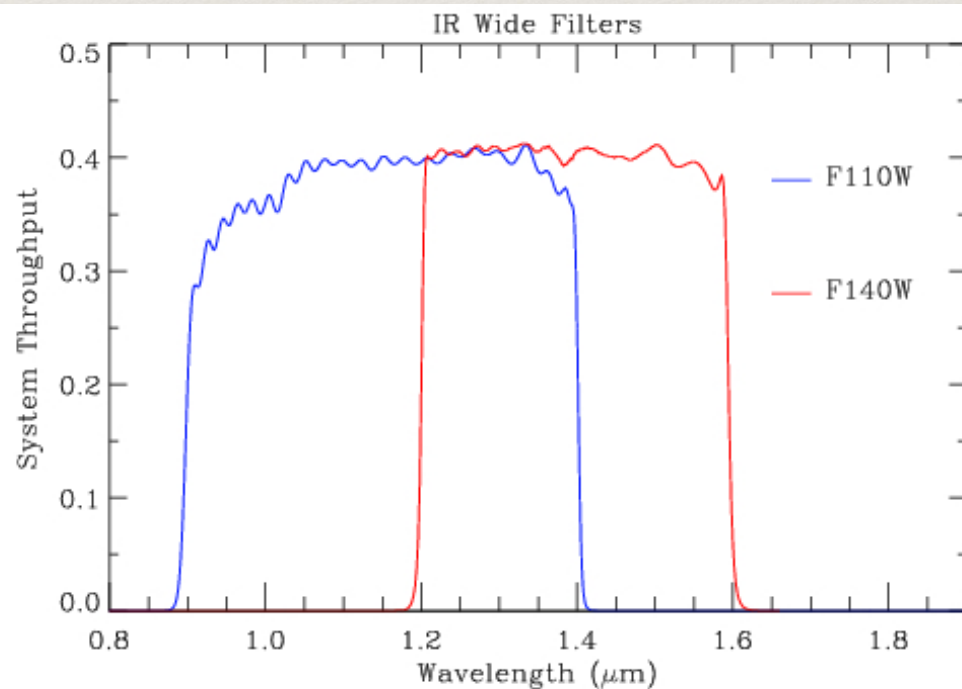
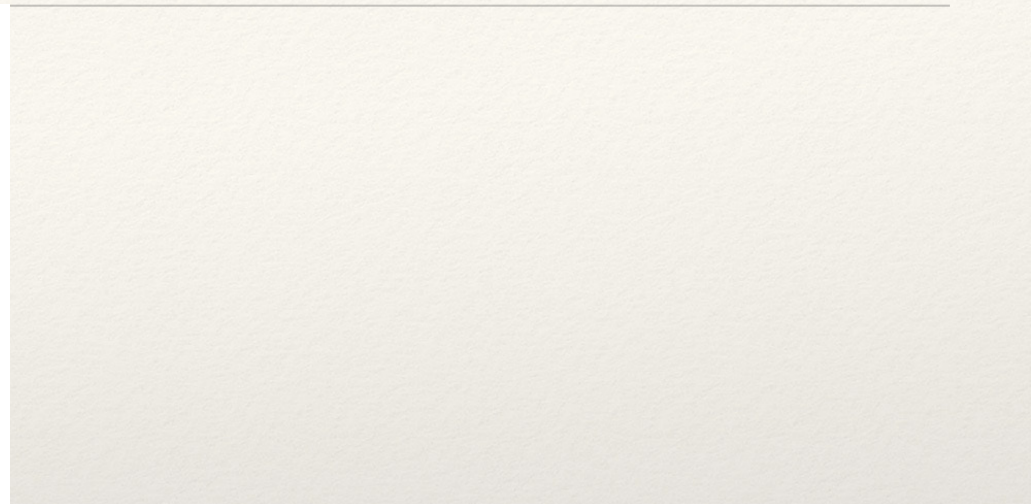
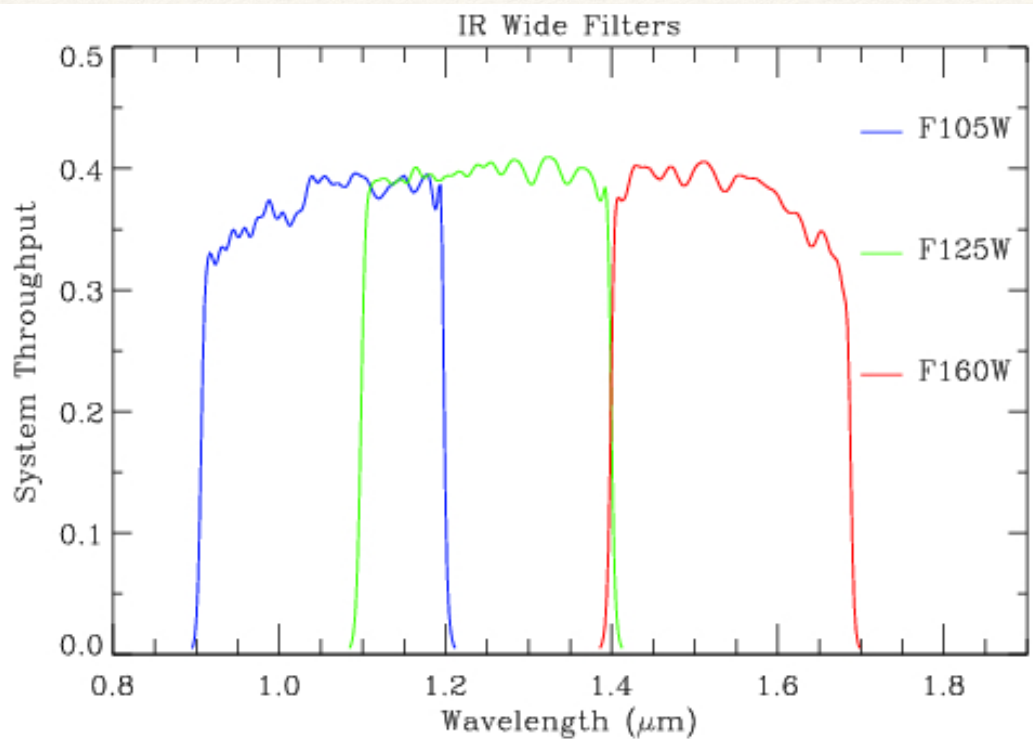
Stromgren v, V, [SII]



410nm,502nm,673nm



# WFC3 for HST

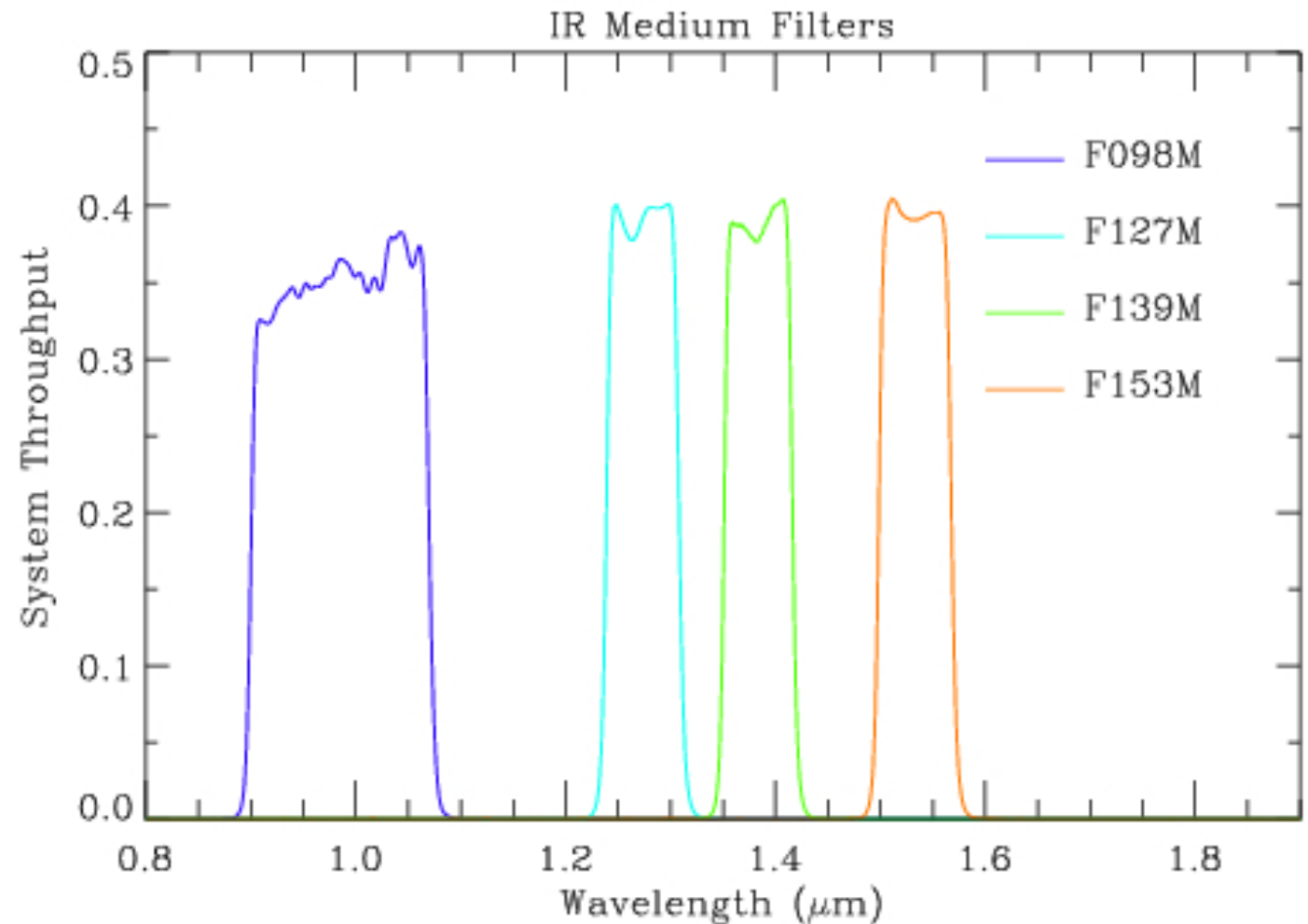


WFC3

Broad band IR

# WFC3 for HST

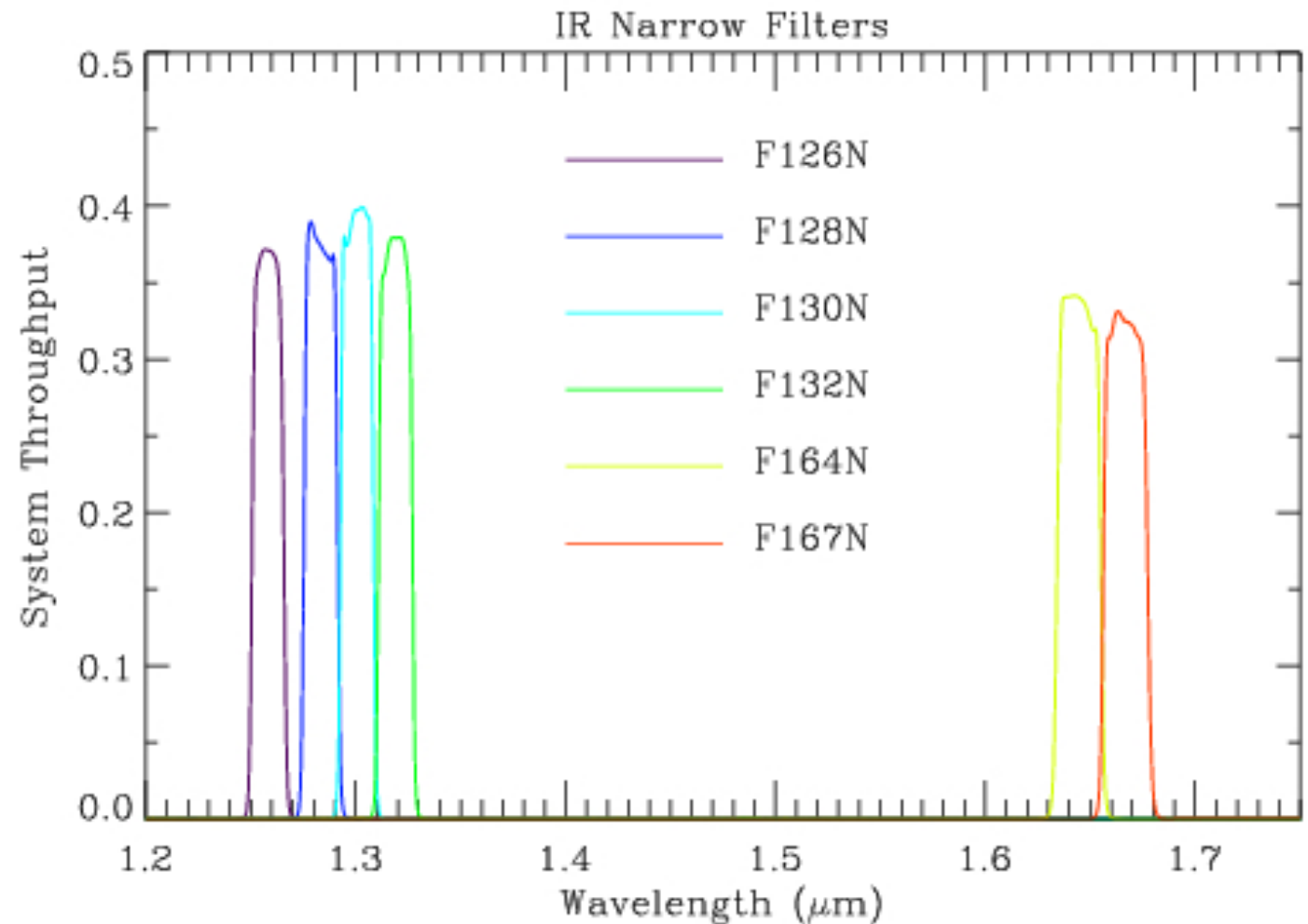
WFC3  
Medium band IR





# WFC3 for HST

WFC3  
Narrow band IR



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# Filter Wheels

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